

Massive neutron stars and Λ -hypernuclei in relativistic mean field models

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Based on relativistic mean field (RMF) models, we study finite Λ -hypernuclei and massive neutron stars. The effective N - N interactions PK1 and TM1 are adopted, while the N - Λ interactions are constrained by reproducing the binding energy of Λ -hyperon at 1s orbit of $^{40}_{\Lambda}\text{Ca}$. It is found that the Λ -meson couplings follow a simple relation, indicating a fixed Λ potential well for symmetric nuclear matter at saturation densities, i.e., around $V_{\Lambda} = -29.786$ MeV. With those interactions, a large mass range of Λ -hypernuclei can be well described. Furthermore, the masses of PSR J1614-2230 and PSR J0348+0432 can be attained adopting the Λ -meson couplings $g_{\sigma\Lambda}/g_{\sigma N} \gtrsim 0.73$, $g_{\omega\Lambda}/g_{\omega N} \gtrsim 0.80$ for PK1 and $g_{\sigma\Lambda}/g_{\sigma N} \gtrsim 0.81$, $g_{\omega\Lambda}/g_{\omega N} \gtrsim 0.90$ for TM1, respectively. This resolves the Hyperon Puzzle without introducing any additional degrees of freedom.

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I. INTRODUCTION

A proper equation of state (EoS) of baryonic matter is crucial to unveil the dynamics of core-collapse supernova [1], neutron star properties [2–4], binary neutron star mergers [5], and heavy-ion collisions [6, 7]. Currently, the properties of nuclear matter near the saturation density ρ_0 are well constrained. However, the composition of matter at higher densities is still an open question with many possibilities [8, 9]. When density is larger than 2-3 ρ_0 , hyperons (Λ , Σ , Ξ , ...) are created via weak reactions to lower the energy of the system. Being the lightest hyperon, Λ^0 will firstly appear due to an attractive potential in nuclear matter [10, 11], while the heavier ones can only appear at larger densities. For neutron star matter, negatively charged hyperons such as Σ^- and Ξ^- may be important since they can neutralize protons. However, recent studies on the quasi-free Σ^- production spectra [12–14] suggests that the Σ^- potential in nuclear matter is repulsive [10]. If so, Σ -hyperons can only appear at much larger densities.

To understand the properties of baryonic matter at high densities with the possible existence of Λ -hyperons, one needs to construct a proper Λ -baryon interaction. Many attempts were made to extract the Λ -baryon interaction, e.g., a phenomenological N - Λ potential was obtained from the scattering data constraint [15]. However, due to the limited experimental data, there are still large ambiguities for the N - Λ interaction in this scheme. Nevertheless, the structures of Λ -hypernuclei provide crucial constraints for the Λ -baryon interaction.

Based on the experimental studies on the single- Λ hypernuclei [16], one can construct the N - Λ interaction via various nuclear structure models, e.g., shell model [17–20], cluster model [21–25], antisymmetrized molecular dynamics [26], quark mean field model [27], relativistic mean field (RMF) model [28–35], Skyrme-Hartree-Fock model [36], and quark-meson coupling model [37–39]. Note that the Λ - Λ interaction is found to be weakly attractive in the measurement of Λ - Λ bond energies of double- Λ hypernuclei, such as $^{6}_{\Lambda\Lambda}\text{He}$ [40]. Promising results for the N - Λ and Λ - Λ potentials were also obtained from lattice QCD simulations [41].

Based on the obtained Λ -baryon interactions, great successes were achieved on understanding the properties of Λ -hypernuclei. However, when applying those interactions to study neutron stars, the results are controversial regarding the observation of pulsars, i.e., the so called Hyperon Puzzle [42]. Due to the appearance of Λ -hyperons at higher densities, the EoS of baryonic matter becomes soft. As a consequence, the predicted mass of a neutron star can not reach $2M_{\odot}$, which is not in accordance with the recently measured masses for PSR J1614-2230 ($1.928 \pm 0.017 M_{\odot}$) [43, 44] and PSR J0348+0432 ($2.01 \pm 0.04 M_{\odot}$) [45].

To solve the Hyperon Puzzle, one needs to introduce additional repulsion so that the EoS becomes stiffer and consequently predicts neutron stars with a maximum mass exceeding $2M_{\odot}$. As categorized in Ref. [46], there are mainly three types of mechanisms that could afford such repulsive forces: (1) the strange vector meson ϕ in RMF models which results in a repulsive interaction between hyperons [47–51]; (2) three-body forces between hyperons and nucleons in the framework of microscopic many-body theories [52–56]; and (3) the deconfinement phase transition that happens prior to the onset den-

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sity of hyperons [57–64]. Note that the above mentioned mechanisms involve additional degrees of freedom which are not very well constrained. In this paper, we readjust the Λ -meson coupling strengths according to the experimental single- Λ binding energies of Λ -hypernuclei in the framework of the RMF model. The new interactions are then used to calculate the mass-radius (M - R) relation of compact stars, which is compared with the observational masses of pulsars. It is shown that the properties of Λ -hypernuclei and massive neutron stars can be well reproduced without introducing any additional degrees of freedom.

The paper is organized as follows. In Sec. II we present the formalism of RMF models for Λ -hypernuclei and neutron stars. The Λ -meson coupling constants are fixed and applied to study the properties of Λ -hypernuclei and neutron stars in Sec. III, and a summary is given in Sec. IV.

II. THEORETICAL FRAMEWORK

A. Lagrangian Density

RMF models have been shown to be suitable for the studies of finite (hyper)nuclei [65–75] as well as baryonic matter [76–82]. The starting point of the meson-exchange RMF model for baryonic matter is the following covariant Lagrangian density

$$\mathcal{L} = \mathcal{L}_N + \mathcal{L}_Y + \mathcal{L}_l. \quad (1)$$

Here \mathcal{L}_N is the standard RMF Lagrangian density for nucleons [65–68] in which the couplings with the isoscalar-scalar σ , isoscalar-vector ω_μ , isovector-vector $\vec{\rho}_\mu$ mesons, and the photon A_μ are included, i.e.,

$$\begin{aligned} \mathcal{L}_N = \sum_{i=n,p} \bar{\psi}_i [i\gamma^\mu \partial_\mu - M_i - g_{\sigma i} \sigma - g_{\omega i} \gamma^\mu \omega_\mu \\ - g_{\rho i} \gamma^\mu \vec{\tau}_i \cdot \vec{\rho}_\mu - e\gamma^\mu A_\mu \frac{1 - \tau_{i,3}}{2}] \psi_i \\ + \frac{1}{2} \partial_\mu \sigma \partial^\mu \sigma - \frac{1}{2} m_\sigma^2 \sigma^2 - \frac{1}{3} g_2 \sigma^3 - \frac{1}{4} g_3 \sigma^4 \\ - \frac{1}{4} \Omega_{\mu\nu} \Omega^{\mu\nu} + \frac{1}{2} m_\omega^2 \omega_\mu \omega^\mu + \frac{1}{4} c_3 (\omega_\mu \omega^\mu)^2 \\ - \frac{1}{4} \vec{R}_{\mu\nu} \cdot \vec{R}^{\mu\nu} + \frac{1}{2} m_\rho^2 \vec{\rho}_\mu \cdot \vec{\rho}^\mu - \frac{1}{4} F_{\mu\nu} F^{\mu\nu}, \quad (2) \end{aligned}$$

where M_i ($i = n, p$) denotes the nucleon mass, $\vec{\tau}_i$ is the isospin with its 3rd component $\tau_{i,3}$, and $m_\sigma(g_{\sigma i})$, $m_\omega(g_{\omega i})$ and $m_\rho(g_{\rho i})$ are the masses (coupling constants) for the σ -, ω -, and ρ -mesons, respectively. Note that g_2 , g_3 , and c_3 are parameters introduced in the nonlinear self-coupling terms. The field tensors of the vector mesons Ω and \vec{R} and photons F are defined as

$$\Omega_{\mu\nu} = \partial_\mu \omega_\nu - \partial_\nu \omega_\mu, \quad (3a)$$

$$\vec{R}_{\mu\nu} = \partial_\mu \vec{\rho}_\nu - \partial_\nu \vec{\rho}_\mu, \quad (3b)$$

$$F_{\mu\nu} = \partial_\mu A_\nu - \partial_\nu A_\mu. \quad (3c)$$

We adopt the arrows to indicate vectors in isospin space.

The Lagrangian density \mathcal{L}_Y represents the contributions from hyperons [35, 73, 74]. Since Λ hyperons are charge neutral with isospin $\vec{\tau} = 0$, only the couplings with σ - and ω -mesons are included, i.e.,

$$\mathcal{L}_Y = \bar{\psi}_\Lambda [i\gamma^\mu \partial_\mu - M_\Lambda - g_{\sigma\Lambda} \sigma - g_{\omega\Lambda} \gamma^\mu \omega_\mu \\ - \frac{f_{\omega\Lambda}}{2M_\Lambda} \sigma^{\mu\nu} \partial_\nu \omega_\mu] \psi_\Lambda, \quad (4)$$

where M_Λ is the mass of Λ hyperon, $g_{\sigma\Lambda}$ and $g_{\omega\Lambda}$ are coupling constants with σ - and ω -mesons, respectively. The last term in \mathcal{L}_Y is the tensor coupling with ω -meson, which is related with the s.p. spin-orbit splitting.

Meanwhile, the Lagrangian density \mathcal{L}_l is for e and μ leptons with

$$\mathcal{L}_l = \sum_{i=e,\mu} \bar{\psi}_i [i\gamma^\mu \partial_\mu + e\gamma^\mu A_\mu - M_i] \psi_i, \quad (5)$$

and M_i ($i = e, \mu$) are the masses.

For a system with time-reversal symmetry, the space-like components of the vector fields ω_μ and $\vec{\rho}_\mu$ vanish, leaving only the time components ω_0 and $\vec{\rho}_0$. Meanwhile, the charge conservation guarantees that only the 3rd component $\rho_{0,3}$ in the isospin space of $\vec{\rho}_0$ survives. Adopting the mean field and no-sea approximations, the single particle (s.p.) Dirac equations for baryons and the Klein-Gordon equations for mesons and photon can be obtained by a variational procedure.

B. Finite Λ -hypernuclei

To investigate finite Λ -hypernuclei, we neglect leptons since their contributions are comparatively small. In the spherical cases, the Dirac spinor for baryons can be expanded as

$$\psi_{\kappa m}(\mathbf{r}) = \frac{1}{r} \begin{pmatrix} iG_{\kappa}(r) \\ F_{\kappa}(r) \boldsymbol{\sigma} \cdot \hat{\mathbf{r}} \end{pmatrix} Y_{jm}^l(\theta, \phi), \quad (6)$$

with $G_{\kappa}(r)/r$ and $F_{\kappa}(r)/r$ being the radial wave functions for the upper and lower components, while $Y_{jm}^l(\theta, \phi)$ is the spinor spherical harmonics. The quantum number κ is defined by the angular momenta (l, j) as $\kappa = (-1)^{j+l+1/2}(j+1/2)$.

Based on the variational method, the Dirac equation for the radial wave functions of baryons ($i = n, p, \Lambda$) is obtained as

$$\begin{pmatrix} V_i + S_i & -\frac{d}{dr} + \frac{\kappa}{r} + T_i \\ \frac{d}{dr} + \frac{\kappa}{r} + T_i & V_i - S_i - 2M_i \end{pmatrix} \begin{pmatrix} G_{\kappa} \\ F_{\kappa} \end{pmatrix} = \varepsilon_{\kappa} \begin{pmatrix} G_{\kappa} \\ F_{\kappa} \end{pmatrix}, \quad (7)$$

with the s.p. energy ε_{κ} and the mean field scalar, vector

and tensor potentials

$$S_i = g_{\sigma i} \sigma, \quad (8a)$$

$$V_i = g_{\omega i} \omega_0 + g_{\rho i} \tau_{i,3} \rho_{0,3} + \frac{1}{2} e (1 - \tau_{i,3}) A_0, \quad (8b)$$

$$T_i = -\frac{f_{\omega i}}{2M_i} \partial_r \omega_0. \quad (8c)$$

Note that the terms related with $\rho_{0,3}$ and A_0 in Eq. (8b) are zero for Λ hyperons, while the tensor potential in Eq. (8c) is zero for nucleons.

The Klein-Gordon equations for mesons and photons are,

$$(\partial^\mu \partial_\mu + m_\phi^2) \phi = S_\phi, \quad (9)$$

with source terms

$$S_\phi = \begin{cases} \sum_{i=n,p,\Lambda} -g_{\sigma i} \rho_{si} - g_2 \sigma^2 - g_3 \sigma^3, & \phi = \sigma; \\ \sum_{i=n,p,\Lambda} g_{\omega i} \rho_{vi} + \frac{f_{\omega \Lambda}}{2M_\Lambda} \partial_k j_{T\Lambda}^{0k} - c_3 \omega_0^3, & \phi = \omega; \\ \sum_{i=n,p} g_{\rho i} \tau_{i,3} \rho_{vi}, & \phi = \rho; \\ e \rho_c, & \phi = A. \end{cases} \quad (10)$$

where ρ_{si} and ρ_{vi} are the scalar and baryon densities for nucleons, ρ_c is the charge density for protons, and $j_{T\Lambda}^{0k}$ is the tensor density for Λ hyperons.

With the radial wave functions, the densities for the baryons in Eq. (10) can be expressed as

$$\rho_{si}(r) = \frac{1}{4\pi r^2} \sum_{k=1}^{A_i} [|G_{ki}(r)|^2 - |F_{ki}(r)|^2], \quad (11a)$$

$$\rho_{vi}(r) = \frac{1}{4\pi r^2} \sum_{k=1}^{A_i} [|G_{ki}(r)|^2 + |F_{ki}(r)|^2], \quad (11b)$$

$$\rho_c(r) = \frac{1}{4\pi r^2} \sum_{k=1}^{A_p} [|G_{kp}(r)|^2 + |F_{kp}(r)|^2], \quad (11c)$$

$$j_{T\Lambda}^0 = \frac{1}{4\pi r^2} \sum_{k=1}^{A_\Lambda} [2G_{k\Lambda}(r)F_{k\Lambda}(r)] \mathbf{n}, \quad (11d)$$

where \mathbf{n} is the angular unit vector. The baryon number A_i ($i = n, p, \Lambda$) can be calculated by integrating the baryon density $\rho_{vi}(r)$ in coordinate space as

$$A_i = \int 4\pi r^2 dr \rho_{vi}(r). \quad (12)$$

For given N - N and N - Λ effective interactions, we solved the Dirac Eq. (7), mean field potentials Eq. (8), Klein-Gordon Eq. (9), and densities Eq. (11) in the RMF model by iteration in coordinate space with a box size of $R = 20$ fm and a step size of 0.05 fm.

C. Neutron Stars

To investigate neutron stars with RMF models, the procedure is similar to that used for finite nuclei. However, a neutron star is comprised of approximately 10^{57}

baryons with the radius being typically around 10 km, which is much larger than a nucleus. When we consider a large number of particles in a large system, the exact solutions of the boundary problem can be replaced by plane waves. Then the meson fields in neutron star matter are determined by

$$m_\sigma^2 \sigma = \sum_{i=n,p,\Lambda} -g_{\sigma i} \rho_{si} - g_2 \sigma^2 - g_3 \sigma^3, \quad (13a)$$

$$m_\omega^2 \omega_0 = \sum_{i=n,p,\Lambda} g_{\omega i} \rho_{vi} - c_3 \omega_0^3, \quad (13b)$$

$$m_\rho^2 \rho_{0,3} = \sum_{i=n,p} g_{\rho i} \tau_{i,3} \rho_{vi}, \quad (13c)$$

with the source currents of baryon i

$$\rho_{vi} = \langle \bar{\psi}_i \gamma^0 \psi_i \rangle = \frac{g_i \nu_i^3}{6\pi^2}, \quad (14a)$$

$$\rho_{si} = \langle \bar{\psi}_i \psi_i \rangle = \frac{g_i m_i^{*3}}{4\pi^2} \left[x_i \sqrt{x_i^2 + 1} - \text{arcsinh}(x_i) \right]. \quad (14b)$$

Here we have defined $x_i \equiv \nu_i/m_i^*$ with ν_i being the Fermi momentum and $g_i = 2$ the degeneracy factor for particle type i . The effective masses of baryons are given as $m_i^* \equiv m_i + g_{\sigma i} \sigma$, while the meson fields σ , ω_0 , and $\rho_{0,3}$ are obtained as their mean values. Note that the charge density is zero in neutron star matter since the local charge neutrality condition should be fulfilled, i.e.,

$$\sum_i q_i \rho_{vi} = 0, \quad (15)$$

with q_i being the charge of particle type $i = (p, e^-, \mu^-)$. Also, the tensor potential T_Λ and density $j_{T\Lambda}^{0k}$ for Λ hyperons vanish in uniform neutron stars.

At zero temperature, the energy density can be determined by

$$E = \sum_i \varepsilon_i(\nu_i, m_i^*) + \frac{1}{2} m_\sigma^2 \sigma^2 + \frac{1}{3} g_2 \sigma^3 + \frac{1}{4} g_3 \sigma^4 + \frac{1}{2} m_\omega^2 \omega_0^2 + \frac{3}{4} c_3 \omega_0^4 + \frac{1}{2} m_\rho^2 \rho_{0,3}^2, \quad (16)$$

with the kinetic energy density of fermion i as

$$\begin{aligned} \varepsilon_i(\nu_i, m_i^*) &= \int_0^{\nu_i} \frac{g_i p^2}{2\pi^2} \sqrt{p^2 + m_i^{*2}} dp \\ &= \frac{g_i m_i^{*4}}{16\pi^2} \left[x_i (2x_i^2 + 1) \sqrt{x_i^2 + 1} - \text{arcsinh}(x_i) \right]. \end{aligned} \quad (17)$$

Note that effective masses of leptons are their own, i.e., $m_{e,\mu}^* = m_{e,\mu}$. The chemical potentials for baryons μ_b ($b = n, p, \Lambda$) and leptons μ_l ($l = e, \mu$) are obtained from

$$\mu_b = g_{\omega b} \omega_0 + g_{\rho b} \tau_{b,3} \rho_{0,3} + \sqrt{\nu_b^2 + m_b^{*2}}, \quad (18a)$$

$$\mu_l = \sqrt{\nu_l^2 + m_l^2}. \quad (18b)$$

Then the pressure is determined by

$$P = \sum_i \mu_i \rho_{vi} - E. \quad (19)$$

To reach the lowest energy, particles will undergo weak reactions until the β -equilibrium condition is satisfied, i.e.,

$$\mu_\Lambda = \mu_n, \quad \mu_e = \mu_n - \mu_p, \quad \mu_\mu = \mu_e. \quad (20)$$

The EoS of neutron star matter is obtained with Eqs. (16) and (19), then the structure of a neutron star is determined by solving the Tolman-Oppenheimer-Volkov (TOV) equation

$$\frac{dP}{dr} = -\frac{GME}{r^2} \frac{(1 + P/E)(1 + 4\pi r^3 P/M)}{1 - 2GM/r}, \quad (21)$$

with the subsidiary condition

$$M(r) = \int_0^r 4\pi E r'^2 dr'. \quad (22)$$

Here the gravitational constant $G = 6.707 \times 10^{-45} \text{ MeV}^{-2}$.

III. RESULTS AND DISCUSSIONS

We carried out extensive calculations to investigate the properties of Λ -hypernuclei and neutron stars based on the effective N - N interactions PK1 [83] and TM1 [84]. Note that the effective interaction PK1 has been widely adopted in our previous studies on the properties of ordinary nuclei and hypernuclei, while the effective interaction TM1 has been commonly used for supernova simulations, i.e., the Shen EoSs [11]. For the N - Λ interactions, the scalar coupling constant $\alpha_{\sigma\Lambda} \equiv g_{\sigma\Lambda}/g_{\sigma N}$ is constrained by reproducing the experimental binding energies of Λ hyperon in the $1s_{1/2}$ state of hypernucleus $^{40}_{\Lambda}\text{Ca}$ ($B_\Lambda = 18.7 \text{ MeV}$) [16], which are 0.618 and 0.620 for the effective interaction PK1 and TM1, respectively; the vector coupling constant $\alpha_{\omega\Lambda} \equiv g_{\omega\Lambda}/g_{\omega N} = 0.666$ is fixed according to the naive quark model [85]; and the tensor coupling constant $f_{\omega\Lambda} = -1.0g_{\omega\Lambda}$ is taken as in Ref. [31]. The masses of Λ^0 , e^- and μ^- are taken from the Particle Data Group [86]. With those N - N and N - Λ interactions, we calculated the mass dependence of the single- Λ binding energies of Λ -hypernuclei and present the results in Fig. 1 along with their experimental values. Note that the theoretical single- Λ binding energies in Fig. 1 are the average values of those for spin up and spin down orbits. It is found that, with the present N - Λ interactions based on the PK1 and TM1 effective N - N interactions, the RMF model can well describe the hypernuclei in a large mass range of $A = 16 \sim 208$, especially for the heavy hypernucleus $^{208}_{\Lambda}\text{Pb}$.

However, the aforementioned parameter sets are not unique. In fact, as long as the depth of Λ mean field potential

$$V_\Lambda \equiv g_{\sigma\Lambda}\sigma + g_{\omega\Lambda}\omega_0 \quad (23)$$

is fixed, the predicted single- Λ binding energies barely vary with $\alpha_{\sigma\Lambda}$ or $\alpha_{\omega\Lambda}$ [35]. This is shown clearly in Fig. 2,

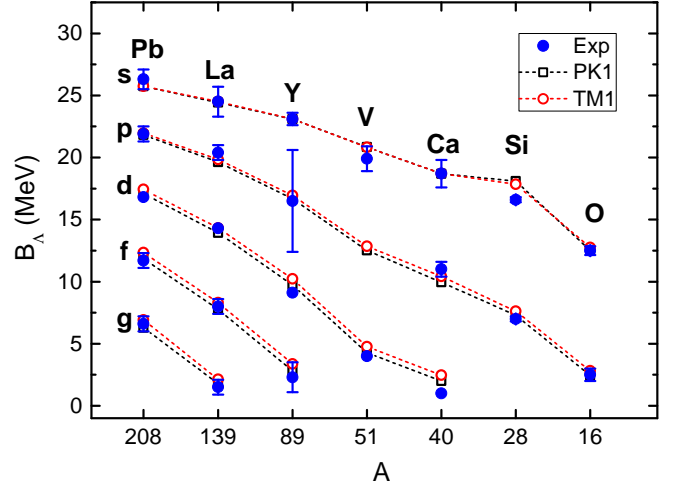


FIG. 1. (Color online) The predicted single- Λ binding energies of Λ -hypernuclei based on the effective interactions PK1 [83] and TM1 [84], which are compared with the experimental data [16].

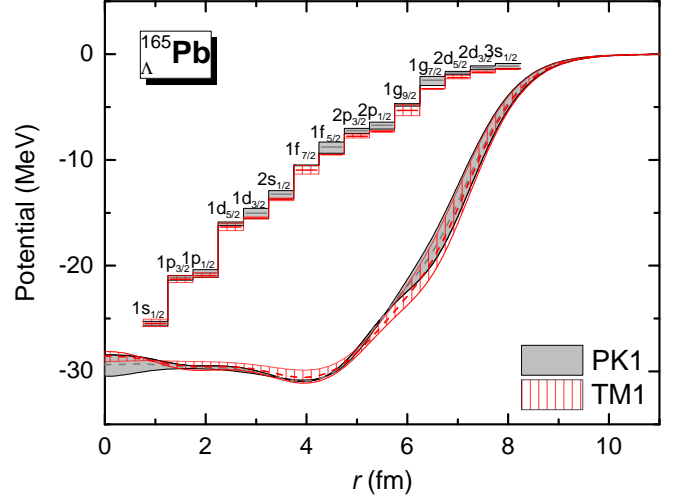


FIG. 2. (Color online) The mean field potential and single-particle levels of Λ in $^{165}_{\Lambda}\text{Pb}$.

TABLE I. The coupling constants $\alpha_{\omega\Lambda}$ and $\alpha_{\sigma\Lambda}$ for N - Λ interactions. The corresponding onset densities ρ_0^Λ (in fm^{-3}) and chemical potentials μ_0^Λ (in MeV) for Λ hyperons are listed here as well.

$\alpha_{\omega\Lambda}$	PK1			TM1		
	$\alpha_{\sigma\Lambda}$	ρ_0^Λ	μ_0^Λ	$\alpha_{\sigma\Lambda}$	ρ_0^Λ	μ_0^Λ
0.60	0.565	0.296	1105.03	0.567	0.308	1104.96
0.65	0.605	0.300	1109.19	0.607	0.312	1109.22
0.70	0.645	0.304	1113.75	0.647	0.317	1113.76
0.75	0.686	0.309	1118.67	0.687	0.323	1118.74
0.80	0.726	0.313	1124.05	0.727	0.328	1124.15
0.85	0.767	0.319	1129.96	0.767	0.334	1130.03
0.90	0.807	0.325	1136.50	0.807	0.341	1136.49
0.95	0.847	0.331	1143.73	0.847	0.348	1143.67
1.00	0.888	0.338	1151.97	0.887	0.357	1151.62

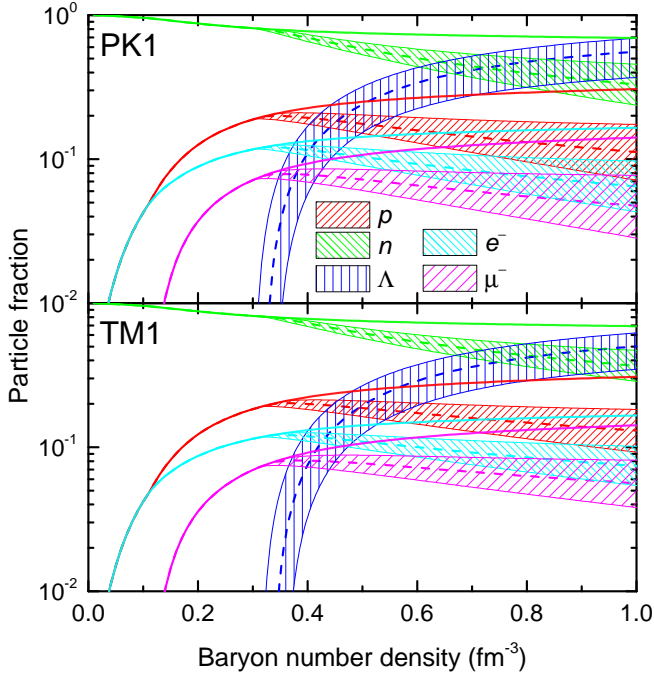


FIG. 3. (Color online) Particle fractions n_i/n of baryons (p, n, Λ) and leptons (e, μ) in neutron star matter. The shaded region are obtained by varying $\alpha_{\omega\Lambda}$ from 0.666 to 1, while the dashed lines in the shadows and solid lines above the shadows correspond to the results obtained with $\alpha_{\omega\Lambda} = 0.85$ and without hyperons, respectively; the same convention is adopted for the following figures.

where the mean field potential and single-particle levels of Λ hyperon in $^{165}_{\Lambda}\text{Pb}$ are presented at various choices of $\alpha_{\sigma\Lambda}$ and $\alpha_{\omega\Lambda}$. Specifically, the shaded region in Fig. 2 are obtained by varying $\alpha_{\omega\Lambda}$ from 0.666 to 1, while the dashed lines represent the central values obtained with $\alpha_{\omega\Lambda} = 0.85$. The σ - Λ couplings are fixed to give $V_{\Lambda} = -29.786$ MeV for symmetric nuclear matter at saturation densities. For various choices of $\alpha_{\omega\Lambda}$, the corresponding values of $\alpha_{\sigma\Lambda}$ that reproduce the binding energies of Λ -hypernuclei are listed in Table I. It is found that the Λ potential varies little for different Λ -meson couplings or nuclear effective interactions. Correspondingly, the single- Λ binding energies are well constrained within 1 MeV for $^{165}_{\Lambda}\text{Pb}$ as indicated in Fig. 2. Similar behaviors are observed for other Λ -hypernuclei as well.

At a given total baryon number density n , the properties of neutron star matter can be obtained by simultaneously fulfilling the conditions of baryon number conservation with $n = \rho_{vn} + \rho_{vp} + \rho_{v\Lambda}$, charge neutrality in Eq. (15), and chemical equilibrium in Eq. (20). As was done for $^{165}_{\Lambda}\text{Pb}$ in Fig. 2, we do not specify the exact values for $\alpha_{\sigma\Lambda}$ and $\alpha_{\omega\Lambda}$, but rather varying $\alpha_{\omega\Lambda}$ from 0.666 to 1 while $\alpha_{\sigma\Lambda}$ is determined by Eq. (23) with $V_{\Lambda} = -29.786$ MeV. The particle number density for each species is determined by Eq. (14a), where the particle fractions are presented in Fig. 3 as functions of the total baryon number density n . As indicated in Table I,

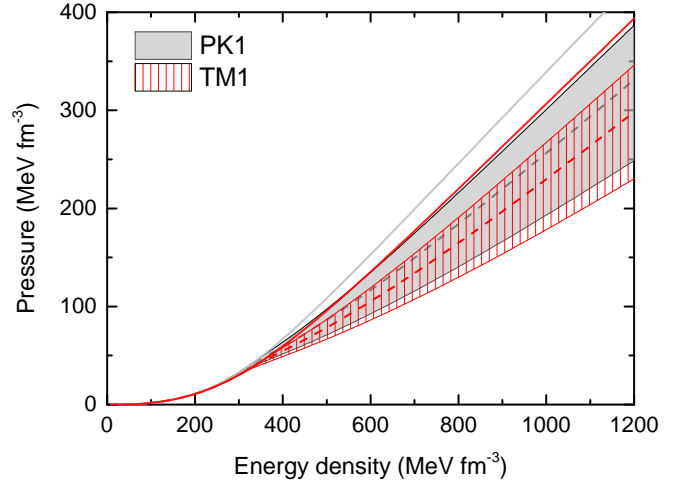


FIG. 4. (Color online) The pressure P of neutron star matter as functions of the energy density E .

varying $\alpha_{\omega\Lambda}$ has minor impact on hyperonic matter at lower densities, where the onset density for Λ increases from 0.30 to 0.36 fm^{-3} as we increase the values of $\alpha_{\omega\Lambda}$. At larger densities, the neutron star matter is dominated by Λ starting at $n = 0.6$ to 1.4 fm^{-3} , depending on the adopted interactions.

The energy density E and pressure P of neutron star matter are obtained with Eqs. (16) and (19) along with the subsidiary equations in Sec. II C. In Fig. 4, we present the pressure of neutron star matter as functions of energy density. Comparing with nuclear matter (solid gray line), it is found that the EoS is softened once Λ -hyperon appears (shaded area) at approximately 300 MeV fm^{-3} . With larger $\alpha_{\omega\Lambda}$, the EoS becomes stiffer due to the increasing repulsive interaction which originated from ω -mesons at larger densities.

Now the structures of neutron stars can be determined by solving the TOV equation in Eq. (21) based on the EoSs displayed in Fig. 4. At the densities below half of the nuclear saturation density ($\sim 0.08 \text{ fm}^{-3}$), we adopt the EoS given in Refs. [88–90] since uniform nuclear matter becomes unstable and a transition to crystalized matter forming the neutron star crust takes place. In Fig. 5, we present the obtained M - R relations of neutron stars with possible existence of hyperonic matter. If we take the commonly-used value of $\alpha_{\omega\Lambda} = 0.666$ according to the naive quark model [85], the obtained maximum neutron star mass is lower than $1.8 M_{\odot}$, in disagreement with the observed masses of PSR J1614-2230 and PSR J0348+0432 ($\sim 2 M_{\odot}$). We note, however, that only a slight adjustment of $\alpha_{\omega\Lambda}$ is necessary to reach the observed masses of these pulsars, specifically $\alpha_{\omega\Lambda} \gtrsim 0.8$ for PK1 and $\alpha_{\omega\Lambda} \gtrsim 0.9$ for TM1. Correspondingly, the σ - Λ couplings should be larger, i.e., $\alpha_{\sigma\Lambda} \gtrsim 0.73$ for PK1 and $\alpha_{\sigma\Lambda} \gtrsim 0.8$ for TM1, to obtain agreement between our predictions and the pulsar observations.

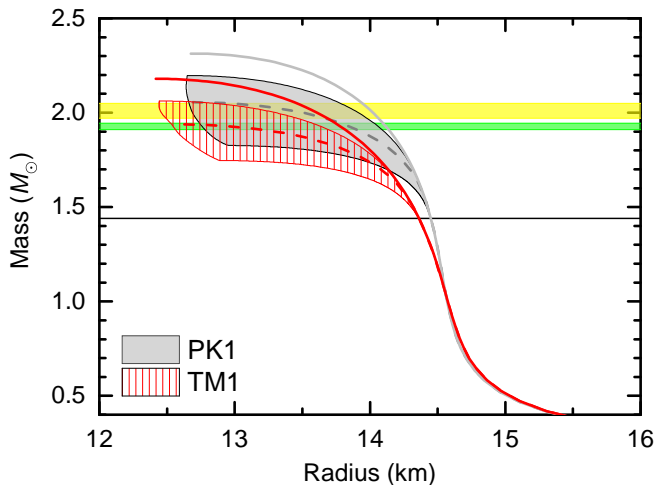


FIG. 5. (Color online) The obtained M - R relations of neutron stars including the possible existence of Λ -hyperons. The masses of pulsars PSR J1614-2230 ($1.928 \pm 0.017 M_\odot$) [43, 44], PSR J0348+0432 ($2.01 \pm 0.04 M_\odot$) [45], and PSR 1913+16 ($1.4398 \pm 0.0002 M_\odot$) [87] are indicated with horizontal bands.

IV. SUMMARY

The possible existence of Λ -hyperons in neutron stars were explored in RMF models, in which we use the effective interactions PK1 and TM1 for nucleons while the Λ -meson couplings were constrained according to the experimental single- Λ binding energies of Λ -hypernuclei. A simple relation between the Λ -meson couplings ($g_{\omega\Lambda}$ and $g_{\sigma\Lambda}$) was obtained which gives a constant value for the depth of Λ potential $V_\Lambda = -29.786$ MeV for symmetric nuclear matter at saturation densities. With these baryon-meson couplings, we studied the properties of neutron star matter including Λ -hyperons and found that the onset densities for Λ lie in the range of 0.30-

0.36 fm^{-3} . The EoSs were softened at densities above which Λ -hyperons appear. However, the EoS is less softened if we adopt larger values of $\alpha_{\omega\Lambda} = g_{\omega\Lambda}/g_{\omega N}$. By solving the TOV equation with new EoSs, we found that the maximum mass of neutron stars can reach $2 M_\odot$ if we use $\alpha_{\omega\Lambda} \gtrsim 0.8$ for PK1 and $\alpha_{\omega\Lambda} \gtrsim 0.9$ for TM1, values just slightly higher than those used in the naive quark model. Thus, we conclude that the values of Λ -meson couplings should be close to those of nucleon-meson couplings so that the single- Λ binding energies agree with measured data and the maximum mass of neutron stars is consistent with the latest observed masses of pulsars.

Furthermore, concerning the recently observed neutron star merger gravitational wave (GW170817) [91] and the possible gravitational wave signal of a post-merger remnant [92], the EoSs and $\alpha_{\omega\Lambda}$ may be even further constrained. Meanwhile, heavier hyperons such as Ξ^- , which possibly play important roles at large densities, have not been considered. The effects of Ξ^- to the properties of neutron star matter, and accordingly to the structure of compact stars, should be considered in our future works.

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