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High resolution X-ray spectroscopy of the stellar wind in Vela X-1 during a flare

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ABSTRACT

Context. We present a ∼130 ks observation of the prototypical wind accreting high-mass X-ray binary Vela X-1 collected with *XMM-Newton* at orbital phases between 0.12 and 0.28. A strong flare took place during the observation allowing us to investigate the reaction of the clumpy stellar wind to the increased X-ray irradiation.

Aims. To examine the wind's reaction to the flare, we perform both time-averaged and time-resolved analyses of the RGS spectrum and examine potential spectral changes.

Methods. We focus on the high-resolution *XMM-Newton* RGS spectra and divide the observation into pre-flare, flare, and post-flare phases. We model the time-averaged and time-resolved spectra with phenomenological components and with the self-consistent photoionization models calculated via CLOUDY and XSTAR in the pre-flare phase, where strong emission lines due to resonant transitions of highly ionized ions are seen.

Results. In the spectra, we find emission lines corresponding to K-shell transitions in highly charged ions of oxygen, neon, magnesium, and silicon as well as radiative recombination continua (RRC) of oxygen. Additionally, there are potential absorption lines of magnesium at a lower ionization stage and features identified as iron L lines. The CLOUDY and XSTAR photoionization models provide contradictory results, pointing towards uncertainties in theory and/or possibly a more complex multi-phase plasma.

Conclusions. We are able to show the existence of a plethora of variable narrow features, including the firm detection of oxygen lines and RRC that RGS enables to observe in this source for the first time. We show that Vela X-1 is an ideal source for future high-resolution missions, such as *XRISM* and *Athena*.

Key words. stars: massive – stars: winds, outflows – X-rays: binaries

1. Introduction

1.1. Vela X-1

Vela X-1 (4U 0900–40) is an eclipsing high mass X-ray binary (HMXB) in which a neutron star (NS) closely orbits a supergiant mass donor with a period of ∼9 days (Hiltner et al. 1972; Forman et al. 1973), accreting from the strong stellar wind and emitting brightly in the X-rays. The system lies relatively close, with a Gaia distance estimate of $2.42^{+0.19}_{-0.17}$ kpc (Bailer-Jones et al. 2018), such that despite its intrinsically moderate X-ray luminosity ($\sim 4 \times 10^{36}$ erg s^{−1}, Nagase et al. 1986) it is among the brightest persistent sources on the X-ray sky and has been well observed ever since its detection (Chodil et al. 1967). It is often considered as “prototype” wind-accreting HMXB, and its system parameters or approximations have been used in various modelling studies, see, e.g., Blondin et al. (1990, 1991), Manousakis et al. (2014) or El Mellah et al. (2019).

The mass donor, HD 77581, has a mass of $\sim 24 M_{\odot}$ and a radius of about $30 R_{\odot}$ (van Kerkwijk et al. 1995; Quaintrell et al. 2003; Rawls et al. 2011). Its mass loss rate has been estimated by different authors over the years – see Kretschmar et al. (2019) for an overview – to be in the range of $\sim 0.5\text{--}2 \times 10^{-6} M_{\odot}$ yr^{−1}, while estimates for the terminal velocity of the wind, v_{∞} , have ranged from 1700 km s^{−1} (Dupree et al. 1980) to values of 600–700 km s^{−1} in more recent studies (Krtička et al. 2012; Giménez-García et al. 2016; Sander et al. 2018).

The X-ray source is a relatively massive NS, with estimates ranging from $\sim 1.8 M_{\odot}$ to $\sim 2.3 M_{\odot}$ (Barziv et al. 2001; Quaintrell et al. 2003; Rawls et al. 2011). It orbits HD 77581 at an average distance of $53.4 R_{\odot}$ or $1.7 R_{\text{HD 77581}}$ on an orbit with low eccentricity ($e = 0.0898 \pm 0.0012$, Bildsten et al. 1997). It is a persistent X-ray pulsar with a pulse period of ~ 283 s, which varies randomly on longer timescales (Deeter et al. 1987; Bildsten et al. 1997). The pulsar has a high magnetic field of a few 10^{12} G that can be directly measured from the detected cyclotron lines (e.g., Fürst et al. 2014, and references therein). Being deeply embed-

ded in the dense stellar wind close to the mass donor, the neutron star shows highly variable accretion and absorption depending on the changing properties and structure of the stellar wind (e.g., Malacaria et al. 2016; Fürst et al. 2010), but can also influence the wind flow through its X-ray radiation (e.g., Blondin et al. 1990; Krtička et al. 2012; Sander et al. 2018).

The eclipse of the X-ray source in Vela X-1 covers $\sim 20\%$ of the orbit duration, (e.g., Falanga et al. 2015). In the literature, orbital phases are sometimes expressed relative to the mid-eclipse time, sometimes relative to the time of mean longitude (T_{90} or $T_{\pi/2}$; used in this work). Kreykenbohm et al. (2008) determined the offset between these two definitions as 0.026 ± 0.005 in orbital phase.

After eclipse egress, the observed absorption (characterized by the equivalent hydrogen column density, N_{H}) usually decreases strongly by more than an order of magnitude, as the neutron star slowly emerges from behind the extended atmosphere or inner wind base of the supergiant star (Nagase et al. 1986; Lewis et al. 1992), often reaching minimal absorption around orbital phase 0.25, where the neutron star is mainly moving towards the observer (Doroshenko et al. 2013). At later phases the observed N_{H} can vary strongly by an order of magnitude (e.g. Orlandini et al. 1998) and shows different peaks and minima in different observations, but always also a marked overall increase towards eclipse ingress (e.g. Nagase et al. 1986; Haberl & White 1990; Malacaria et al. 2016). The evolution of N_{H} is usually attributed to the changing sight-line through the complex accretion geometry, which may include an accretion wake, a photoionization wake and a tidal stream (e.g., Blondin et al. 1990; Kaper et al. 1994; van Loon et al. 2001; see Grinberg et al. 2017 for a visualisation and further references). On top of the changes along the orbit, there is significant irregular variation both in absorption as in intrinsic X-ray flux, including bright flares and states of very low flux – “off states” (e.g., Haberl & White 1990; Kreykenbohm et al. 2008; Odaka et al. 2013; Martínez-Núñez et al. 2014; Sidoli et al. 2015).

Early X-ray line studies in Vela X-1 have mainly been focused on the prominent iron line complex, see Becker et al. (1978, *OSO 8*) and Ohashi et al. (1984, *Tenma*), for example. Sato et al. (1986, *Tenma*) also detected a broad Fe $\text{K}\alpha$ emission line, but they suspected that fluorescent $\text{K}\alpha$ lines of Si, S, Ar, Ca, and Ni as well as the Fe $\text{K}\beta$ line also contribute to the overall line intensity. Various further elements have been reported from spectra taken during, or close to eclipse (Nagase et al. 1994; Sako et al. 1999; Schulz et al. 2002), including recombination lines and radiative recombination continua (RRC), as well as fluorescence lines. The variety of ionization states strongly indicates the presence of an inhomogeneous wind with optically thick, less ionized matter coexisting with warm photoionized plasma, e.g., in the form of clumps intrinsic to the wind or in the accretion wake (Grinberg et al. 2017).

Goldstein et al. (2004) and Watanabe et al. (2006) each analysed *Chandra*-HETGS data from three orbital phases: in eclipse, around phase 0.25 and around phase 0.5. In eclipse and in the heavily absorbed observation at phase 0.5 a large number of emission lines were identified, while around phase 0.25 the spectrum was dominated by the continuum and the line content was not discussed further. Grinberg et al. (2017) revisited the *Chandra*-HETGS observation around phase 0.25, creating separate spectra for time intervals of low and high spectral hardness, i.e., performing an absorption-resolved analysis. They detected line features from high and low ionization species of Si, Mg, and Ne, as well as strongly variable absorption. Again, this implies co-existence of cool and hot gas phases, that can either be ex-

plained by an intrinsically clumpy stellar wind or a highly variable, structured accretion flow close to the compact object or a combination of both effects.

1.2. X-ray studies of winds in HMXBs and modelling efforts

Studying stellar winds in HMXBs is not only interesting for research on accretion in X-ray binaries, they are also crucial to understand the stellar winds from massive stars. The diagnostics of the stellar wind structure and properties obtained from X-ray observations in HMXBs are complementary to diagnostics at other wavelengths and for single stars and may be used to break degeneracies in derived models of stellar winds from massive stars in general. Martínez-Núñez et al. (2017) provide an extensive review of theory and observations of inhomogeneous winds in HMXBs as well as in isolated massive stars. As described in this review, various authors have used a wide range of approaches to estimate wind parameters and indications for large and small-scale structures. These include studies of absorption variations by intervening material (e.g., Grinberg et al. 2015; Miškovičová et al. 2016), the variability of the intrinsic X-ray continuum as observed in flares and off-states (e.g., Bozzo et al. 2013; Martínez-Núñez et al. 2014; Pradhan et al. 2019), or fluorescence line studies (e.g., Boroson et al. 2003; Wojdowski et al. 2003). It can be difficult, though, to directly compare results of particular studies, since different effects on the actual observables may be hard to disentangle and, in addition, underlying assumptions frequently vary between studies.

The huge range in length scales to be treated – from binary system scales of $> 10^8$ km to accretion column scales of < 1 km – as well as the many, interconnected physical effects to be taken into account, ideally dynamically self-consistent, have so far precluded a detailed, fully realistic model of a HXMB accreting X-ray pulsar. Existing model efforts for HMXBs treat only certain aspects of the system in detail, while others are **prohibited** or hugely simplified. This limits the ability to test observations directly against predictions, but some comparisons are both possible and have pushed the field forward (e.g., Manousakis & Walter 2015; Grinberg et al. 2017).

Previous studies of clumpy winds in HMXBs with RGS include 4U1700-37 by van der Meer et al. (2004) and 4U0114+65 by Sanjurjo-Ferrín et al. (2017). The former work did not manage to obtain much information about the highly absorbed source from the RGS spectra due to a low count rate despite strong flares. Sanjurjo-Ferrín et al. (2017), on the other hand, were able to detect several emission lines with RGS below 1 keV: C v $\text{Ly}\alpha$, N vi He-like triplet, O vii $\text{Ly}\alpha$, Fe xvii, and Fe xviii. No flares were detected in this observation.

In the following, we use *XMM*-RGS observations of Vela X-1 at orbital phase ~ 0.25 during a strong flare in order to probe the stellar wind and accretion structure properties in the system and, in particular, the wind’s response to changes in X-ray irradiation. We particularly note that this is the first analysis of Vela X-1 with RGS, enabling the coverage of lower energies than previous studies with *Chandra*-HETGS. This work makes use of previous results of the *EPIC*-pn analysis of the same observation in Martínez-Núñez et al. (2014) which is shortly introduced, together with the general discussion of the data used, in Sec. 2. We present our spectral analysis, including a blind line search and photoionization modelling in Sec. 3, discuss our finding in Sec. 4 that also includes an outlook on future prospects of observations with *XRISM* and *Athena* and finish with a summary in Sect. 5.

2. Data

2.1. Observations

Vela X-1 was observed on May 25–26, 2006 (MJD 53 880.439 – MJD 53 881.877; ObsID 0406430201) using the European Photon Imaging Camera (EPIC) and the Reflection Grating Spectrometers (RGS, den Herder et al. 2001) on-board the *XMM-Newton* observatory (Jansen et al. 2001) during its complete revolution 1183. The observations lasted for \sim 130 ks at orbital phases $\phi_{\text{orb}} = 0.12$ –0.28 according to the ephemeris in Kreykenbohm et al. (2008), with phase zero defined as $\phi_{90} = 0$. When observing, the EPIC-Metal Oxide Semi-conductor (EPIC-MOS1, Turner et al. 2001) and RGS 1 cameras were disabled to guarantee the maximum telemetry for the EPIC-pn camera (Strüder et al. 2001). However, on May 26 the RGS 1 camera was switched on since the telemetry bandwidth would still allow to use it (Fig. 1). The EPIC-pn data from this observation were studied in Martínez-Núñez et al. (2014).

2.2. Previous analysis of simultaneous EPIC-pn data

Our analysis builds on the previous results obtained with EPIC-pn, in particular on the knowledge won about the evolution of the continuum emission during the observation. In this section, we give a short summary of those key findings in Martínez-Núñez et al. (2014) that drive our analysis and interpretation.

The authors subdivided the observation into individual spectra, each of 1.1 ks length to trace the evolution of the continuum and absorption throughout the flare. They saw a strong flare during the observation, with an increase in observed EPIC-pn count-rate by a factor of $\gtrsim 10$. They found that this increase is driven both by an increase in the intrinsic luminosity of the source and a decrease in absorption. The phenomenological analysis of the unabsorbed flux and hydrogen column density evolution throughout the flare suggested two physical sources were visible: the neutron star emission and the scattered emission by the wind. The authors propose that the origin of the flare might be associated with accretion of a dense wind clump onto the neutron star.

Martínez-Núñez et al. (2014) also noted a decrease in the column density in two of the absorbing components that they interpret as being located in the vicinity of the neutron star. This was partially explained by accretion of the dense clump and subsequent reduction of the amount of matter in the vicinity of the neutron star. Another reason given was the decrease in the amount of stellar wind material along the line of sight with orbital phase as the neutron star moves along the orbit after eclipse egress.

The EPIC-pn camera covers a wider energy range (~ 0.5 –10 keV) than the RGS camera (~ 0.33 –2.5 keV), however with a significantly lower resolving power. The EPIC-pn spectra in Martínez-Núñez et al. (2014) show hints of spectral line features at energies below ~ 2 keV, and in fact, those have also been observed and identified in *Chandra*-HETG observations of Vela X-1 (Goldstein et al. 2004; Watanabe et al. 2006; Grinberg et al. 2017), although none of these works addressed strong flares as seen in our observation. In this work, we aim to characterize these narrow features and their changes throughout the flare using RGS with its higher resolving power.

2.3. Data reduction

Data reduction was performed in *XMM-Newton* Science Analysis System (SAS) v. 17.0.0 (Gabriel et al. 2004) using the most up-to-date calibration files available in October 2018. We processed the data with the standard SAS tool `rgsproc` to generate event files, spectra and response matrices. We also produced exposure-corrected and background-subtracted RGS 1 and RGS 2 light curves with the `rgs1ccorr` task. From that, we defined the good time intervals (GTIs) and generated three sets of spectra with the `tabgtigen` task (pre-flare, flare and post-flare phases as shown in Fig. 1). The flare phase begins after 76.1 ks ($\phi_{\text{orb}} = 0.215$), and the post-flare phase starts after 99.6 ks ($\phi_{\text{orb}} = 0.245$). The spectra were then rebinned with an `ftools` command `ftgrouppha` with an optimal binning scheme from Kaastra & Bleeker (2016).

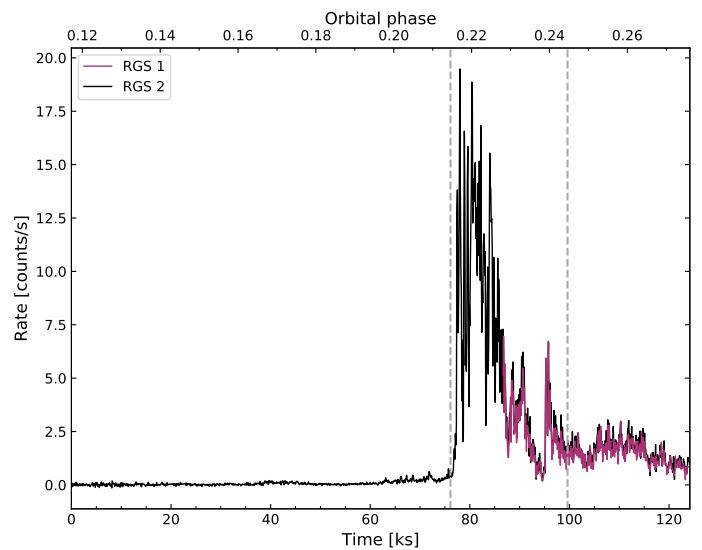


Fig. 1: RGS light curves of Vela X-1 during the observation discussed in this paper. Time on the abscissa are from the beginning of the observations (MJD 53 880.439). The dashed lines show the pre-, during and post-flare phases. We use the time of mean longitude of 90 degrees, T_{90} , from Kreykenbohm et al. (2008) as phase zero.

The publicly available Observation Data Files (ODFs) suffer from time jumps resulting in an error message when running the `rgs1ccorr` task as it cannot handle these jumps properly. The correct version of the ODFs without the time jumps was received directly from the *XMM-Newton* Science Operation Centre. In addition, due to failures of individual RGS CCDs¹, there are gaps in the spectra at ~ 0.9 –1.2 keV for RGS 1, and ~ 0.5 –0.6 keV for RGS 2.

As the background rate for both RGS cameras showed that the maximum recommended value of 2 counts/s was not exceeded, we did not remove intervals of high particle background. This explains why our total lightcurves (MJD 53 880.439 to MJD 53 881.878) are longer than the ones presented in Martínez-Núñez et al. (2014) (MJD 53 880.613 to MJD 53 881.768), who applied such a correction.

¹ https://xmm-tools.cosmos.esa.int/external/xmm_user_support/documentation/uhb/rgs.html

3. Spectral analysis and results

The main goal of the spectroscopic analysis discussed in this paper is the diagnostics of the plasma in the stellar wind. For this purpose, a robust method is needed to identify and characterize narrow-band features, i.e., emission and absorption lines and radiative recombination continua (RRC). We followed a two-tier approach. We first selected a list of narrow-band feature candidates through a blind search, following the method discussed in Tombesi et al. (2010). The candidate line list was later used as an input to a formal forward-folding fitting procedure. For the latter step we used XSPEC (Arnaud 1996) and, in particular, PyXspec. Throughout this work we use the Minuit2 `migrad` minimization method and, due to the low numbers of counts per spectral bin, the C-statistic (Cash 1979). We have analyzed:

- in the pre-flare phase: only the RGS 2 spectrum as no RGS 1 data are available;
- in the flare phase: the spectra of RGS 1 and RGS 2 separately and assuming independent spectral parameters for the two cameras as RGS 1 covers a much shorter part of the flare than RGS 2;
- in the post-flare phase: RGS 1 and RGS 2 spectra simultaneously assuming the same spectral parameters as RGS 1 and RGS 2 cover the same time interval; cf Fig. 1. To take into account possible differences in cross-calibration of the two cameras, we have included a cross-normalization constant, the best-fitting parameter of which is consistent with 1 (0.98 ± 0.04), as expected (de Vries et al. 2015).

In the remainder of this section, we introduce the continuum model in Sect. 3.1, explain the method for the blind line search in Sect. 3.2, present the detected narrow-band spectral features in Sect. 3.3, the direct fits of the R and G parameters in Sect. 3.4, and conclude with photoionization modelling of the data in Sect. 3.5.

3.1. Continuum model

In the analysis, we implemented a continuum model that builds on the previous model of Martínez-Núñez et al. (2014) and provides an appropriate fit for the spectra throughout the whole observation, i.e., in the pre-, during and post-flare phases. Our overall model, including the narrow-band components, is described by:

$$F(E) = \text{tbabs}_1 \times \text{cabs} \times \text{powerlaw}_1 + \text{tbabs}_2 \times \text{powerlaw}_2 + \sum_{i=1}^n G_i + \sum_{i=1}^m RRC_i \quad (1)$$

i.e., two absorbed power laws (defined by normalization and the photon index, Γ) with an addition of n Gaussian absorption or emission line profiles (defined by line energy, width and flux) and m RRC (redge in xspec, defined by threshold energy, plasma temperature and overall normalization).

Unlike Martínez-Núñez et al. (2014), who used *XMM-Newton* EPIC-pn spectra, our model includes only two power law continuum components instead of three. Their third component mainly contributes above ~ 2 keV and is thus not required here given the spectral coverage of RGS. The lack of coverage at higher energies and thus of a good constraint on the underlying, non-absorbed continuum, in general prevents a direct comparison between the parameters of the continuum components

in Martínez-Núñez et al. (2014) and in this paper. However, the choice of the continuum modelling does not have a significant impact on the main results discussed here: the continua in Eq. 1 only represent an empirical description, and are not used to draw any physical conclusions about the state of the plasma in the stellar wind that is based on the narrow-band features that only require a good local continuum description (e.g. Goldstein et al. 2004; Grinberg et al. 2017; van den Eijnden et al. 2019).

The two power law models of our continuum have distinct normalizations and photon indices. To account for the effects occurring in the material with a given column density, N_H , that is situated between the X-ray source and the detector, the power law component is modified by absorption, modelled by two different contributions. We used `tbabs` (Wilms et al. 2000) with abundances from Wilms et al. (2000)² to model the overall absorption and follow Martínez-Núñez et al. (2014) in further using optically-thin non-relativistic Compton scattering, `cabs`. In the highly time-resolved analysis of Martínez-Núñez et al. (2014), where short-term strong increases of absorption, especially at the very beginning of the observation, could be resolved, the inclusion of the model was crucial. Since RGS spectra cannot be analyzed on timescales of a few ks, we do not expect `cabs` to have a significant contribution as it should only become significant at average column densities of $\gtrsim 10^{24}$ cm⁻². Nevertheless, since this component was present in the original continuum model in (Martínez-Núñez et al. 2014), we keep it for historical and consistency reasons. The values of the continuum parameters we obtain are summarized in Tab. 1.

In Fig. 2 we compare the broadband instrument-response corrected RGS 2 spectra during the three phases. The quiet pre-flare phase shows a flat, very low continuum contribution with many strong narrow emission features visible already by eye. The flare causes the continuum to increase drastically as observed by both RGS cameras, and its contribution remains high compared to the pre-flare phase, even in the post-flare phase. Martínez-Núñez et al. (2014) saw a similar behavior in their analysis of the simultaneous EPIC-pn spectra. They, in particular, could show that the difference between the observed pre- and post-flare continuum is driven both by changes in the intrinsic flux of the power law and by changes in absorption, while the underlying spectral shape remained fairly stable.

3.2. Blind line search

To search for faint and narrow features in the time-averaged³ and time-resolved spectra, we opted for the algorithm introduced by Tombesi et al. (2010) which generates energy-intensity contour plots. The blind line search allows to detect spectral features that otherwise might be too weak to be detected with the naked eye.

In this approach, a Gaussian profile with centroid energy E and intensity I is added to the best-fit continuum (cf. Eq. 1) and moved through the given energy range in steps to search for spectral line-like features. We used four different spectral regions corresponding to parts of the spectrum where resonant transitions of highly-ionized species of a given element are expected. The corresponding energy ranges and numbers of steps in energy are listed in Tab. 2. Since we are interested in narrow features, a piece-wise approach to spectral analysis is appropriate and has been often used in the past in other analyses (e.g. Grinberg et al. 2017; van den Eijnden et al. 2019). Our steps in energy

² Note that modern version of `tbabs` automatically uses the cross-sections of Verner et al. (1996).

³ Integrated over the entire observation.

Table 1: Continuum model parameters and corresponding uncertainties. The error bars were calculated for a 90% confidence interval and two interesting parameters.

Phase	$N_{\text{H},1}$ [10^{22} cm^{-2}]	Γ_1	N_1 [ph/keV/cm ² /s]	$N_{\text{H},2}$ [10^{22} cm^{-2}]	Γ_2	N_2 [ph/keV/cm ² /s]	
Time-averaged	$0.52^{+0.04}_{-0.22}$	$-0.4^{+1.8}_{-1.6}$	$(1.2^{+0.1}_{-0.6}) \times 10^{-2}$	$2.54^{+0.05}_{-0.15}$	$1.95^{+0.07}_{-0.24}$	$0.16^{+0.01}_{-0.05}$	RGS 2
Pre-flare	–	–	–	< 0.09	$-1.7^{+0.3}_{-0.2}$	$(4.9^{+0.9}_{-0.4}) \times 10^{-4}$	RGS 2
Flare	$3.3^{+0.1}_{-0.6}$	$2.6^{+0.1}_{-0.7}$	$0.80^{+0.07}_{-0.34}$	$0.9^{+0.4}_{-0.3}$	$1.2^{+1.5}_{-2.4}$	$(4.2^{+6.9}_{-2.2}) \times 10^{-2}$	RGS 1
	$0.9^{+0.2}_{-1.1}$	$1.0^{+0.7}_{-0.5}$	$0.11^{+0.05}_{-0.02}$	$3.27^{+0.06}_{-0.43}$	$2.5^{+1.2}_{-0.5}$	$1.26^{+0.08}_{-0.44}$	RGS 2
Post-flare	$0.9^{+1.1}_{-0.2}$	$-1.1^{+4.0}_{-1.3}$	$(8.9^{+11.8}_{-4.5}) \times 10^{-3}$	$2.1^{+5.0}_{-0.2}$	$1.8^{+5.7}_{-1.4}$	$0.2^{+0.2}_{-0.2}$	RGS 1 & RGS 2

Notes. $N_{\text{H},1}$ refers to the hydrogen column density of both cabs and tbabs₁. The cross-normalization constant used for the simultaneous fit of the post-flare RGS 1 and RGS 2 spectra is $0.98^{+0.04}_{-0.04}$.

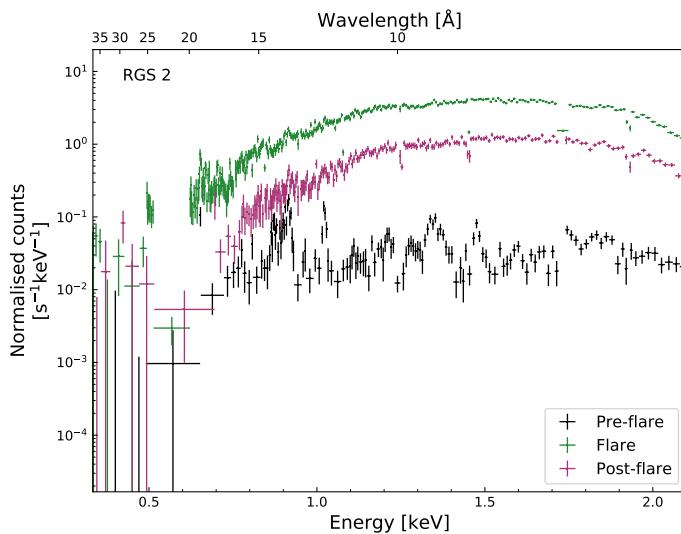


Fig. 2: Re-binned, folded, and time-resolved RGS 2 count spectra (crosses).

correspond to $\sim 1 \text{ eV}^4$. The intensity of the line was varied in 25 steps in the range of -10^{-4} to $10^{-4} \text{ ph s}^{-1} \text{ cm}^{-2}$, thus covering both emission and absorption lines. The difference in goodness-of-fit statistics between the models with and without the Gaussian component were registered at each grid, allowing to build a goodness-of-fit parameter map over the energy vs. intensity plane corresponding to the 68%, 90% and 99% confidence levels for one interesting parameter⁵ (see Fig. 3 in Tombesi et al. (2010) for illustration). From this map, candidate lines were drawn as those corresponding to a confidence level larger than 99% for two interesting parameters.

We then identified the features and modelled them with Gaussian components in case of emission and absorption lines

⁴ The sensitivity of the RGS instruments allows to resolve two lines separated by $\Delta E \sim 1\text{--}2 \text{ eV}$ at $\sim 0.4\text{--}0.6 \text{ keV}$ and $\Delta E \sim 7\text{--}8 \text{ eV}$ at $\sim 1.4 \text{ keV}$ (https://xmm-tools.cosmos.esa.int/external/xmm_user_support/documentation/uhb/rgsresolve.html).

⁵ See Lampton et al. (1976) for definition of interesting parameters. Generally, the number of interesting parameters affects the range of their confidence intervals.

Table 2: Regions used for line searches and number of steps used in each region.

Main expected element	Energy range [eV]	N_{steps}
O	300–900	667
Ne	800–1400	667
Mg	1200–1700	556
Si	1600–2100	667

and the redge model in case of RRC. Several weak spectral features suggested by the blind line search were subsequently removed from our final model as they were judged to be instrumental or otherwise artificial. In particular, we disregarded a possible line at 732.01 eV detected in a feature-rich region: the region around $\sim 730 \text{ eV}$ is complex, with two of our continuum components intersecting and Ovn RRC present (at 0.739 keV; Lotz 1968). The apparent emission feature at 515.20 eV in the RGS 1 flare spectrum on the other hand, is coincident with a cool pixel at 514.43 eV⁶.

3.3. Narrow-band spectral features

3.3.1. N, O, Ne, Mg, and Si lines

All Gaussian components in our model were assumed to be unresolved (i.e., width = 0) since leaving the intrinsic width free did not yield an improvement in the quality of the fit, at least at the 90% confidence level for two interesting parameters in the time-averaged RGS 2 spectrum. This is consistent with previous results at different orbital phases with *Chandra*-HETG where most lines were unresolved even at HETGs higher resolution (Goldstein et al. 2004; Grinberg et al. 2017). Likewise, we verified that the centroid energy of the lines with an unambiguous identification are consistent with the laboratory energy of the corresponding atomic transition. The calculated line velocities appeared to be consistent with zero which is an expected behavior given our expectation of velocities of a few hundred km/s (e.g., Grinberg et al. 2017) and the resolution of RGS. These measurements were therefore not included in Tabs. 3 and 4.

⁶ https://xmm-tools.cosmos.esa.int/external/xmm_user_support/documentation/uhb/rgsmultipoint.html

Spectra and best fitting models in regions around the detected lines are shown in Fig. 3 to Fig. 5. We performed detailed spectral fits only in the regions where our blind line search did yield significant line detections during any of the phases. In particular, this means that we modelled the 300–650 eV, 600–800 eV, 800–1600 eV, and 1600–2100 eV ranges in all three phases (pre-flare, flare and post-flare) for RGS 1 and 300–550 eV, 600–800 eV, 800–1600 eV, and 1600–2100 eV ranges in all three phases for RGS 2 (cf. Tab. 2).

Ideally, each of the spectral lines and RRC in our model should have been absorbed by a separate tbabs component to avoid any assumptions about what part of absorption affects these line-like features; however, as shown in Eq. 1, they were assumed to be unabsorbed. Due to the relatively narrow energy ranges above and the independence of the line-like features from each other, except for the Ne IX and Mg XI triplet complexes as discussed below, we would not have been able to constrain the tbabs components, thus, no ISM absorption was implemented (cf. the CLOUDY model in Eq. 3 in Sec. 3.5.2).

In summary, we have detected the lines identified as the following transitions: N VII Ly α , O VII He- α , O VIII Ly α , O VII He- β , O VII He- γ , Ne IX He- α (f, i, r), Ne X Ly α , Ne X Ly β , Mg VI Ka, Mg VII Ka, Mg XI He- α (f, r), Mg XII Ly α , and Si XIII He- α (r) as well as O VIII RRC. The best-fit parameters for all the Gaussian profiles in Eq. 1 detected at a confidence level higher than 90% (two interesting parameters) are shown in Tab. 3 for the time-averaged spectrum and Tab. 4 for the time-resolved spectra, together with the most likely identification. The detected RRC are listed in Tab. 5 and discussed in Sect. 3.3.2. Finally, the three transitions with ambiguous identifications are presented in Tab. 6 and discussed in Sect. 3.3.3.

In the case of the observed He-like triplets (Ne IX and Mg XI), we fitted them as a combination of three Gaussian components, for which relative intensity was expressed as a function of the R and G parameters, as described in Sec. 3.4. The line energies of each component in the Ne IX triplet were tied to the Ne X Ly α line using transition reference energies as published in the literature (see Tab. 4 for references) under the assumption that these lines belong to the same dynamic region and, thus, are affected by the same Doppler shift. For the Mg XI triplet, the intercombination and forbidden lines were tied to the resonance line in a similar manner.

3.3.2. Radiative recombination continua

At 865.25 eV, we observed a narrow emission feature in the pre-flare phase spectrum corresponding to the O VII RRC, see Tab. 5. We modelled this feature with the *redge* model in XSPEC. At 1362.77 eV, and thus coincident with the Mg XI-triplet (Tab. 4), we expect a further RRC feature, namely the Ne X RRC. This feature has been previously seen in the eclipse spectra of Vela X-1 using *Chandra*-HETG (Schulz et al. 2002). Similarly, the feature is seen by *Chandra*-HETG in the highly absorbed Vela X-1 spectra at $\phi_{\text{orb}} \approx 0.75$ (Amato et al., in prep.). Although the blind line search does not explicitly find this feature among the Mg triplet lines given the resolution of RGS, we included it in our overall model to assess its possible contribution. We again used the *redge* model. The RRC features appeared to be narrow with $kT_e = 13.9^{+17.6}_{-6.8}$ eV for O VIII and $kT_e < 12.6$ eV for Ne X (Tab. 5), and both can be significantly detected only during the pre-flare phase.

3.3.3. Possible presence of Fe L transitions

In addition to the already discussed features in Tab. 5, the blind line search suggests further weak lines around 1050 eV, 1067 eV, and 1115 eV (Tab. 6). Previous *Chandra*-HETG observations of Vela X-1 during eclipse show hints of emission features around these energies, identified with Fe L transitions, although the lines were not explicitly modelled (Schulz et al. 2002).

The energies of the lines at 1050 eV and 1115 eV are, within the fitted centroid uncertainties, compatible with some Fe L transitions, namely the Fe XXII B13 line at 11.770 Å (1053.392 eV), the Fe XXIII Be2 line at 11.736 Å (1056.443 eV), and the Fe XXIV Li3 line at 11.176 Å (1109.379 eV), as measured with EBIT (Brown et al. 2002). All three are the strongest observed transitions in the respective ionization stages for a collisionally ionized plasma; we note however, that at least a part of the emission we observe is likely from a photoionized component (Sect. 3.5). The line at 1067 eV, on the other hand, is at best compatible with weak unresolved lines in B-like Fe XXII and F-like Fe XVIII while two slightly stronger lines in Be-like Fe XXIII (Be3 at 11.702 Å or 1059.513 eV and Be4 at 11.458 Å or 1082.075 eV) and one line in Li-like Fe XXIV (Li1 11.432 Å or 1084.536 eV) are just outside of the confidence interval.

For all three features in Tab. 6, the fitted line centroids have comparatively large uncertainties, the improvement of the fit statistics is somewhat marginal, and the fitted intensities are basically consistent with zero. If these lines were associated with L-shell transitions in Fe ions, we would expect to see a larger number of these features due to the richness of L-shell transitions in Fe. It is, in principle, possible that these three features are the strongest lines, with others remaining undetected because of the low signal-to-noise ratio. But only the feature at 1050 eV is constrained in all three flare phases, the other two lines only appear during flare and post-flare, respectively; and the lines at 1050 eV and 1067 eV are observed in absorption, while the feature at 1115 eV is in emission. For these reasons, while we cannot rule out the presence of Fe L lines, the data do not support a firm detection either.

3.4. R and G diagnostic parameters

The ratios of the normalization of the forbidden (f), intercombination (i), and resonance (r) lines in He-like triplets can be used to probe temperature, density, and ionizing processes in astrophysical plasmas (Gabriel & Jordan 1969; Porquet & Dubau 2000). They are defined as $R = f/i$ and $G = (f+i)/r$, where the R parameter depends primarily on the electron density, n_e , while the G parameter is density-independent but shows an electron temperature, T_e , dependence instead.

To estimate these parameters, we created a user-defined model in XSPEC containing a sum of three Gaussian functions with three parameters: line energy, width, and intensity. The line intensity of the intercombination and resonance lines were expressed as a factor containing R and/or G according to the equations above, while the normalization of the model corresponded to the line intensity of the forbidden line. This approach allows us to directly fit for R and G .

In Fig. 6 we show 2-D contour plots of the R and G parameters for the detected triplets in the RGS phase-resolved spectra discussed in this paper. The corresponding values of n_e and T_e presented in Tab. 7 were taken from Porquet & Dubau (2000) who have calculated the ratios of the diagnostic parameters in fully or partially photoionized plasmas.

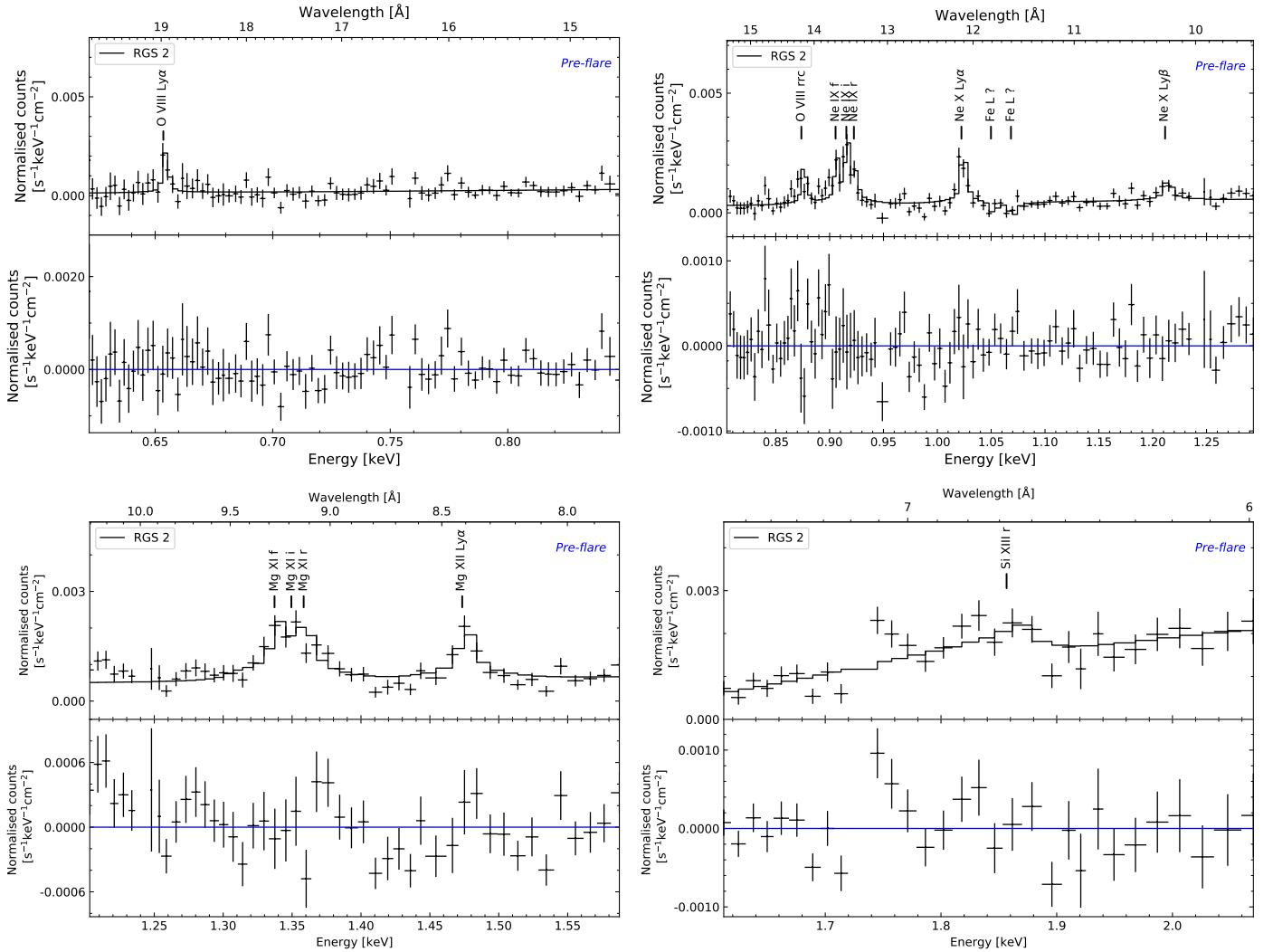


Fig. 3: *Upper panel*: RGS 2 spectrum of the pre-flare phase in four energy bands (crosses), and best-fit model containing Gaussian components and RRS (solid line). *Lower panel*: residuals in units of data-model.

The main results of the analysis can be summarized as follows:

- The R and G diagnostic parameters for the Ne IX-triplet in the pre-flare phase are constrained at $1 \lesssim R \lesssim 3.5$ and $0.5 \lesssim G \lesssim 2.5$ (here and hereafter we refer to the range within the 90% confidence interval).
- The Mg XI-triplet was only detected in the pre-flare phase. The contour plot in the top right panel in Fig. 6 gives $0.5 \lesssim R \lesssim 4$, while for the G parameter the range is basically unconstrained ($G \lesssim 7$).
- For the Ne IX-triplet in the flare phase, there is a weak correlation between the diagnostic parameters. The confidence interval for the R parameter is rather wide: $0.5 \lesssim R \lesssim 6$, while the G parameter reaches $G \lesssim 3.5$ resembling the pre-flare phase.
- In the post-flare phase, we obtained $1.5 \lesssim R \lesssim 6.5$ and $G \lesssim 5$ for the Ne IX-triplet, see the bottom right panel in Fig. 6.

Unsurprisingly, the corresponding ranges of n_e and T_e in Tab. 7 are rather wide, especially in the flare and post flare phases. Given the pre-flare results, the electron number density varies between $\sim 10^{10}$ and 10^{12} cm^{-3} , while the range of the electron temperature is between $\sim 10^5$ and 10^6 K .

In addition, we note that these values should be treated with caution: the results in Porquet & Dubau (2000) do not account for a strong UV radiation field which affects level populations in He-like triplets, i. e., UV pumping. In Sect. 4.3.1, we show that UV pumping is indeed important in the case of Vela X-1.

3.5. Photoionization models

The features detected in the preceding sections, especially in the case of the clear emission lines we observed in the pre-flare spectrum, are indicative of a photoionized medium. We test this assumption through the direct modelling of the pre-flare spectrum with the photoionization model CLOUDY (Ferland et al. 2017).

3.5.1. Input for photoionization models

We approximated the illuminating continuum, required for the simulation of the photoionized wind medium, as a sum of two components, namely the emission from the neutron star and the emission from HD 77581. A similar approach was employed previously in Grinberg et al. (2017). HD 77581 is dominating the UV continuum and is described by a blackbody, assuming

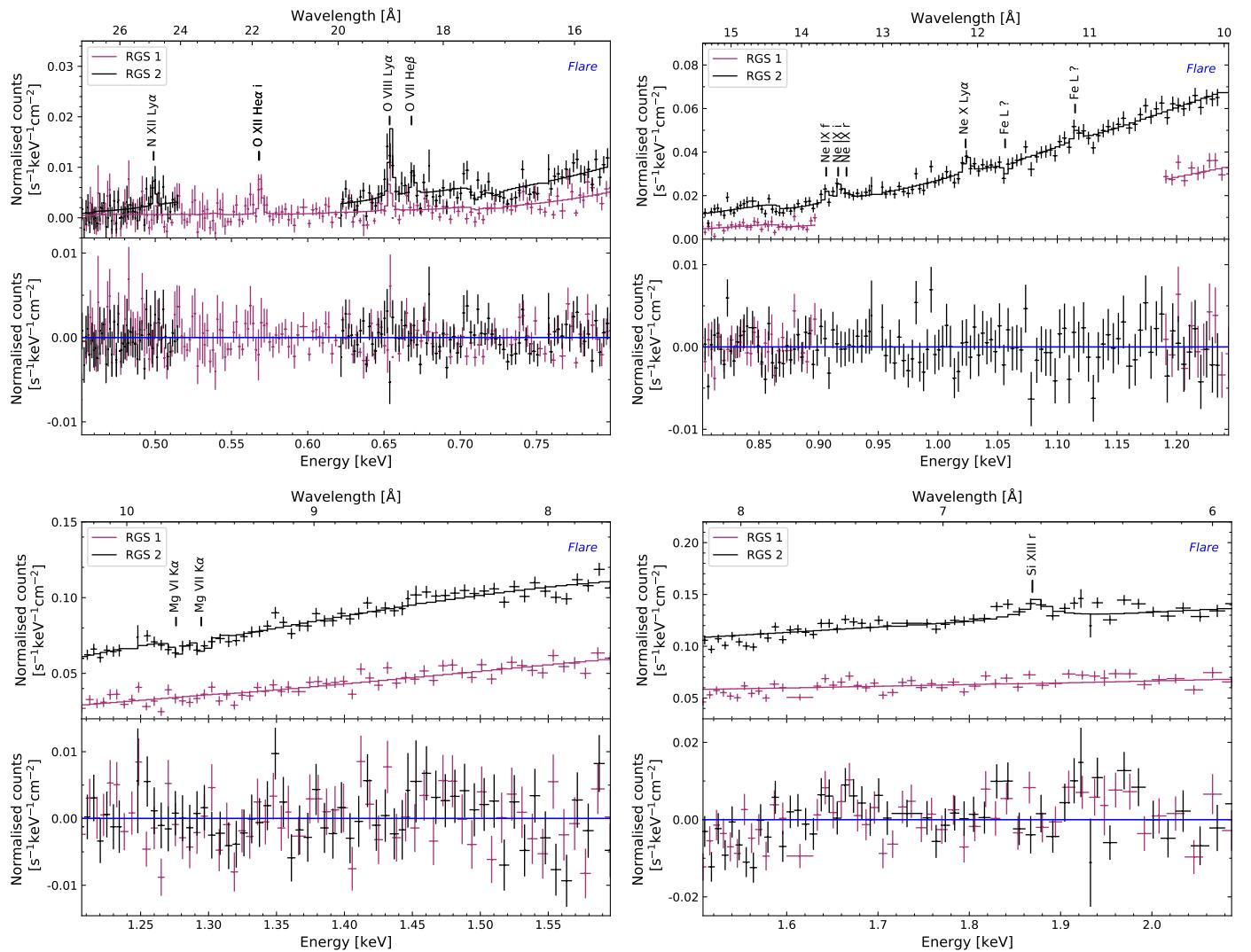


Fig. 4: *Upper panel*: RGS 1 and RGS 2 spectrum of the flare phase in four energy bands (crosses), and best-fit model containing Gaussian components (solid line). *Lower panel*: residuals in units of data-model. The flux difference between RGS 1 and RGS 2 are attributed to the different coverage of the flare, as shown in Fig. 1.

a distance of 2 kpc, stellar radius of $31 R_{\odot}$ and a temperature of 20 000 K.

The limited bandpass of the RGS does not allow us to constrain the broadband continuum of the neutron star. However, previous work by Martínez-Núñez et al. (2014) characterizes the behavior up to 10 keV, in particular the X-ray flux of the source in the 0.6–10 keV range, during our observation. The broadband spectral shape has been best measured with other X-ray telescopes, e.g., with *NuSTAR* by Fürst et al. (2014), who find that a Fermi-Dirac cut-off (FDCUT; Tanaka 1986)

$$F(E) = AE^{-\Gamma} \left(1 + \exp\left(\frac{E - E_{\text{cut}}}{E_{\text{fold}}}\right)\right)^{-1} \quad (2)$$

with A the overall normalization, Γ the photon index, E_{cut} the cut-off energy, and E_{fold} the folding energy, described the observations best. As in Grinberg et al. (2017), we assumed $\Gamma = 0.99$, $E_{\text{cut}} = 20.8$ keV, and $E_{\text{fold}} = 11.9$ keV. We then scaled the normalization of the FDCUT such that the flux corresponded to typical fluxes found by Martínez-Núñez et al. (2014) in the respective time periods. In particular, for the pre-flare phase

a representative value for the unabsorbed 0.6–10 keV flux is $\sim 3.2 \times 10^{-9}$ erg s^{-1} cm^{-2} .

We further estimated the number density of the stellar wind in the vicinity of the neutron star. To do so, we used the mass loss rate, binary separation distance and radius of HD 77581 as described in Sec. 1, while the wind velocity was assumed to be 700 km s^{-1} , i.e. the terminal velocity as measured by (Giménez-García et al. 2016). Assuming a spherically symmetric propagation of the wind and mass-conservation, we estimated the mass density at the location of the neutron star to $\sim 9 \times 10^{-13}$ kg cm^{-3} and the corresponding number density to $\sim 3 \times 10^8$ cm^{-3} . To account for possible wind clumping, we then calculated the photoionization models for a range of densities around this value.

3.5.2. CLOUDY

We generated a set of photoionized emission models in CLOUDY 17.00 (last described in Ferland et al. 2017) for two free parameters: the electron number density, n_e , and the ionization parameter, ξ . The electron number density varied between $5.5 \leq \log n_e/cm^{-3} \leq 11.5$ in 17 steps, while the ionization param-

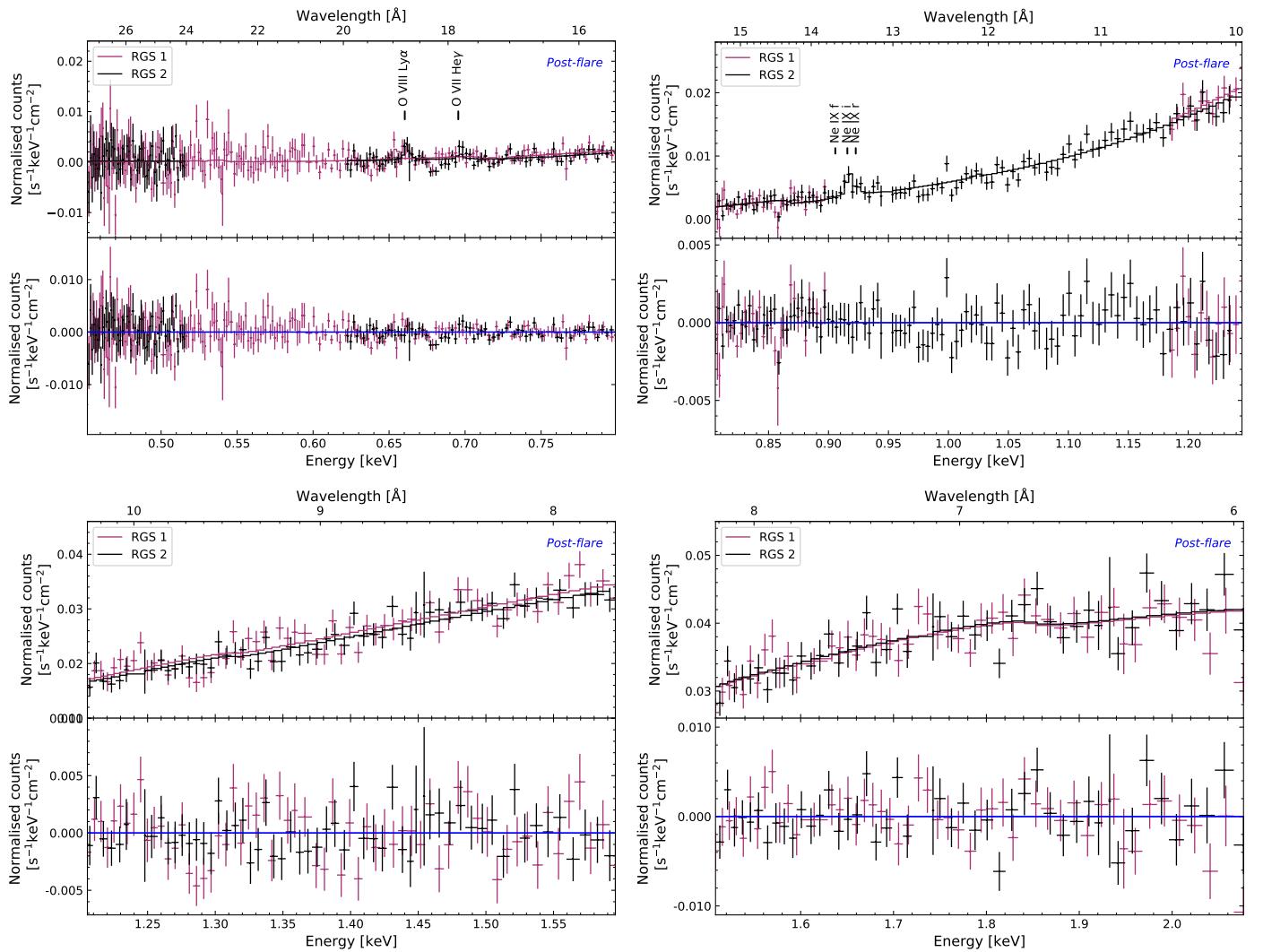


Fig. 5: *Upper panel:* RGS 1 and RGS 2 spectrum of the post-flare phase in four energy bands (crosses), and best-fit model containing Gaussian components (solid line). *Lower panel:* residuals in units of data-model.

parameter varied in 10 steps, such that $0 \leq \log \xi / \text{erg} \cdot \text{cm} \cdot \text{s}^{-1} \leq 4$. Additionally, the turbulent velocity of 10 km s^{-1} was taken from Sander et al. (2018), and the hydrogen column density was fixed at 10^{23} cm^{-2} , adopting the default CLOUDY solar abundances for a slab geometry.

To fit the pre-flare RGS 2 spectrum by interpolating between the model grids, we adopted a model defined as:

$$F(E) = \text{tbabs}_1 \times \text{powerlaw}_1 + \text{tbabs}_2 \times \sum_{i=1}^n \text{CLOUDY}_i \quad (3)$$

i.e., an absorbed power law (defined by normalization and the photon index, Γ) and n absorbed additive CLOUDY models (defined by the electron number density, n_e , the ionization parameter, ξ , and overall normalization).

Unlike the model in the phenomenological analysis shown in Eq. 1, where none of the line-like features were absorbed by tbabs, the n CLOUDY models are absorbed here. The difference lies in the fit being performed for the entire energy range for CLOUDY and only two CLOUDY parameters ($\log n_e$ and $\log \xi$) determining the line-like spectral features over the entire energy range; thus, it was possible to constrain the tbabs component. The overall model was defined as:

Panels (a) and (b) in Fig. 7, show the fitted model with one CLOUDY component and the corresponding residuals, respectively.

Since several lines were not fitted perfectly, e.g., Ne IX (i), Mg XII Ly α , and Si lines at energies above ~ 1.7 keV, we added another CLOUDY component with a lower $\log \xi$ value such that both CLOUDY components were modified by the same tbabs model, see Eq. 3. Indeed, this improved the quality of the fit which can be seen from the residuals in panel (d) in Fig. 7 and from the C-statistic values in Tab. 8: $\Delta C / \Delta \nu = 62/4$. The first component is significantly more dominant as shown in panel (e) in Fig. 7, while the contribution of the low ionization component is mostly important at higher energies in the Si region.

4. Discussion

4.1. Models for wind accretion in HMXBs

Mass accretion in wind-fed HMXBs occurs through accretion of line-driven stellar winds, in which the compact object is embedded. Additionally, there is a possibility of a persistent disk-like structure to be formed around Vela X-1 despite the Roche lobe not being filled by the donor star (El Mellah et al. 2019) – we

Table 3: Detected absorption or emission lines (unresolved) in the time-averaged RGS 1 and RGS 2 spectra.

Ion	E_{ref} [eV]	Line energy [eV]	Normalization [10^{-5} ph cm $^{-2}$ s $^{-1}$]	ΔC	Instrument
O VII He- α (i)	568.7 ^a	569.4 $^{+0.5}_{-1.9}$	2.2 $^{+1.3}_{-1.4}$	10.4	RGS 1
O VIII Ly α	654.12 ^b	653.8 $^{+0.5}_{-1.1}$	2.1 $^{+1.2}_{-1.1}$	18.0	RGS 1
O VIII Ly α	654.12 ^b	653.4 $^{+0.5}_{-0.8}$	1.6 $^{+0.5}_{-0.4}$	38.6	RGS 2
Ne IX (f)	905.74 ^b	904.8 $^{+1.3}_{-2.2}$	1.8 $^{+0.9}_{-0.7}$	28.2	RGS 2
Ne IX (i)	915.58 ^b	916.2 $^{+0.5}_{-1.4}$	3.7 $^{+1.0}_{-0.9}$	96.5	RGS 2
Ne X Ly α	1022.60 ^c	1022.9 $^{+2.0}_{-1.5}$	3.4 $^{+1.2}_{-1.2}$	42.9	RGS 2
Fe L xxiv (?)	1109.379 ^d	1114.2 $^{+4.6}_{-4.0}$	2.6 $^{+1.6}_{-1.4}$	15.1	RGS 2
Mg XI (f/i/r)	1332.11/1344.33/1353.26 ^b	1348.1 $^{+8.3}_{-12.3}$	3.5 $^{+2.9}_{-2.9}$	31.3	RGS 2
Mg VI K α	1276.81 ^e	1275.4 $^{+7.6}_{-10.7}$	-3.7 $^{+2.7}_{-2.0}$	6.3	RGS 2
Mg VII K α	1289.95 ^e	1293.1 $^{+2.9}_{-2.7}$	-6.1 $^{+2.5}_{-2.4}$	10.7	RGS 2

Notes. The statistical significance of the lines is denoted by ΔC , i.e., the difference of the model C-statistic value with and without the line. The error bars were calculated for a 90% confidence interval and two interesting parameters. The line at 1114.24 eV is presumably identified as a Fe L xxiv line, hence, the question mark (see Sec. 3.3.3 for discussion).

References. ^(a) Grant et al. (1980) ^(b) Drake (1988) ^(c) Erickson (1977) ^(d) Brown et al. (2002) ^(e) From Behar & Netzer (2002); value for Mg VII K α averages over the two strong transitions listed there.

Table 5: Detected RRC features in the RGS 2 spectra. The values in the flare and post-flare phase are unconstraining, and therefore, omitted.

		Pre-flare			
Ion	E_{ref} [eV]	E_{obs} [eV]	kT_{e} [eV]	N [ph cm $^{-2}$ s $^{-1}$]	ΔC
O VIII	871.41 ^a	868.0 $^{+3.3}_{-7.3}$	13.9 $^{+17.6}_{-6.8}$	2.4 $^{+1.6}_{-0.9}$	45.3
Ne X	1362.20 ^a	1367.8 $^{+9.3}_{-6.8}$	< 12.6	1.91 $^{+1.1}_{-1.0}$	18.1

Notes. E_{ref} denotes reference threshold energy, E_{obs} is detected threshold energy, kT_{e} is plasma temperature, N is normalization, and ΔC is the difference in the C-statistic with and without the feature in question. The error bars were calculated for a 90% confidence interval and two interesting parameters.

References. ^(a) Garcia & Mack (1965)

Table 7: Statistical ranges for the R and G parameters given the 90% confidence interval (red) in Fig. 6.

Phase	Ion	R	G	n_{e} [cm $^{-3}$]	T_{e} [K]	kT_{e} [eV]
Pre-flare	Ne IX	$1 \lesssim R \lesssim 3.5$	$0.5 \lesssim G \lesssim 2.5$	$10^{11} \lesssim n_{\text{e}} \lesssim 2 \times 10^{12}$	$10^6 \lesssim T_{\text{e}} \lesssim 6 \times 10^6$	$86 \lesssim kT_{\text{e}} \lesssim 517$
Pre-flare	Mg XI	$0.5 \lesssim R \lesssim 4$	$G \lesssim 7$	$10^9 \lesssim n_{\text{e}} \lesssim 5 \times 10^{13}$	$2 \times 10^5 \lesssim T_{\text{e}} \lesssim 8 \times 10^6$	$43 \lesssim kT_{\text{e}} \lesssim 689$
Flare	Ne IX	$0.5 \lesssim R \lesssim 6$	$G \lesssim 3.5$	$10^8 \lesssim n_{\text{e}} \lesssim 2 \times 10^{12}$	$10^6 \lesssim T_{\text{e}} \lesssim 6 \times 10^6$	$86 \lesssim kT_{\text{e}} \lesssim 517$
Post-flare	Ne IX	$1.5 \lesssim R \lesssim 6.5$	$G \lesssim 5$	$10^8 \lesssim n_{\text{e}} \lesssim 10^{12}$	$10^5 \lesssim T_{\text{e}} \lesssim 6 \times 10^6$	$9 \lesssim kT_{\text{e}} \lesssim 517$

Notes. n_{e} denotes electron number density and T_{e} stands for electron temperature. These values were taken from Porquet & Dubau (2000).

note, however, that the high magnetic field of the neutron star means that the disk will not extend close to the compact object and an X-ray emitting disk has indeed never been observed in Vela X-1.

The simplest and widely used wind accretion model described by the Bondi–Hoyle–Lyttleton (BHL) theory (see Edgar 2004 for review) assumes a smooth wind flow to calculate the

mass accretion rate onto the compact object. However, line-driven winds of hot stars are known to be subject to instabilities due to velocity perturbations (e.g., Owocki et al. 1988; Feldmeier et al. 1997; Dessart & Owocki 2005). This leads to strong shocks which in turn create dense gas ‘clumps’ (or shells in 1-D models) surrounded by regions of significantly lower densities (see the review by Martínez-Núñez et al. 2017). Accretion of

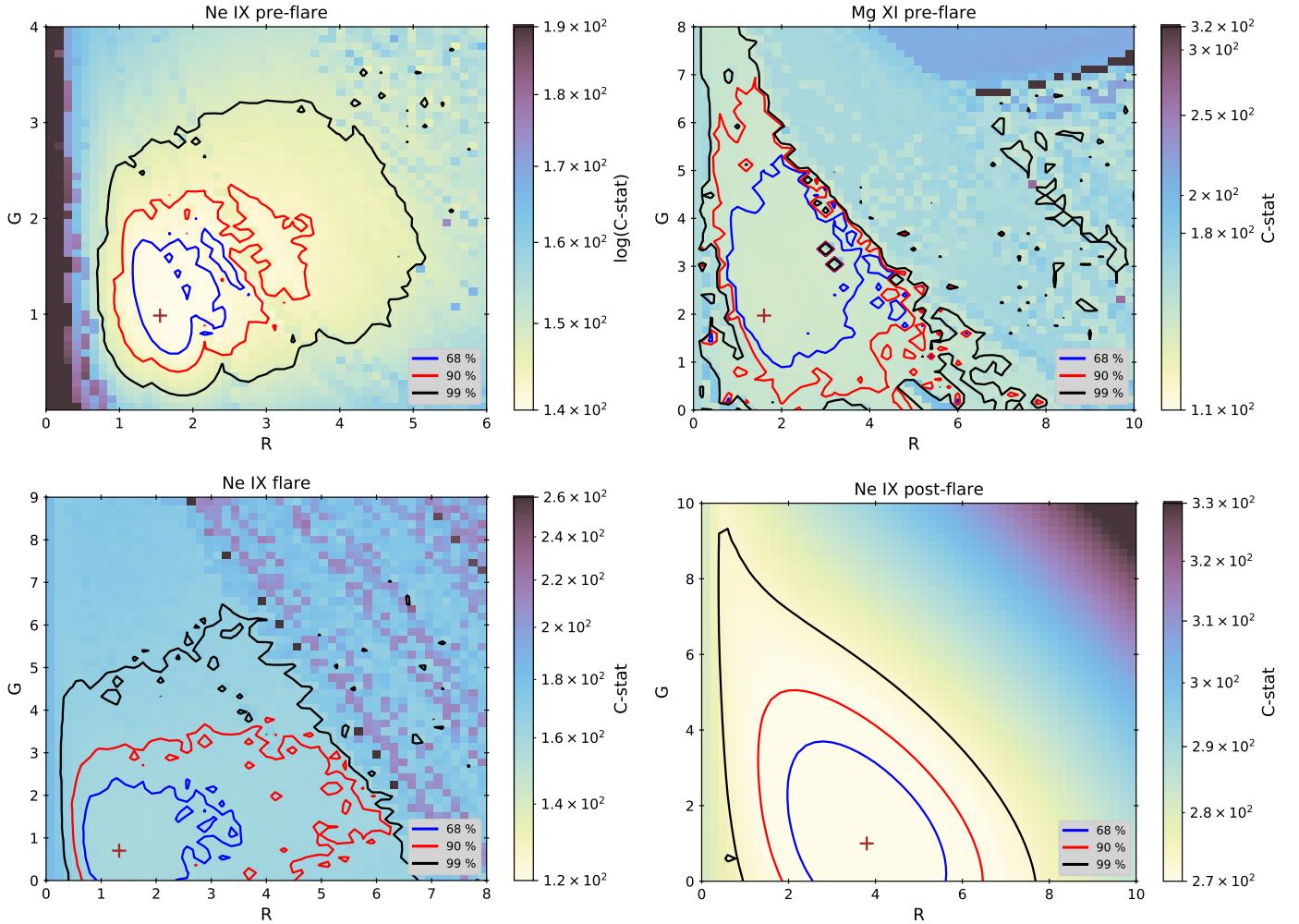


Fig. 6: Iso-C-stat contours in the R vs. G parameter space (50-by-50 grid) for the RGS 2 spectrum. The black, red, and blue contours correspond to 68%, 90%, and 99% confidence intervals, respectively. The crosses denote best-fit values. *Top left*: Ne IX triplet in the pre-flare phase. *Top right*: Mg XI triplet in the pre-flare phase. *Bottom left*: Ne IX triplet in the flare phase. *Bottom right*: Ne IX triplet in the post-flare phase.

clumpy winds has been shown to be less efficient than that of homogeneous winds (El Mellah et al. 2018).

The presence of structured winds is well-established in wind-accreting HMXBs. However, the exact physical properties of these structures are under debate as well as the details of the accretion flow (Martínez-Núñez et al. 2017). Some numerical models (e.g. Blondin et al. 1990; Manousakis & Walter 2015) have described the overall flow, including orbital effects and the impact of the X-ray ionization feedback. Others (e.g. Ducci et al. 2009) have focused on the accretion of individual structures, including also interactions at the magnetosphere and possible changes in the accretion regime (Bozzo et al. 2016). Recently, El Mellah et al. (2018) presented 3D hydrodynamic simulations of the wind in the vicinity of the accretor tracing the inhomogeneous flow over several spatial orders of magnitude, down to the neutron star magnetosphere. However, currently there is no single model treating all the relevant physics across the full range of scales. Thus our measurements cannot yet be tested against a predictive model of the whole accretion process. But they can be used to inform future model developments.

4.2. Discussion of individual feature detections

Before a detailed discussion of the individual line feature detection we emphasize that in the time-resolved spectra the detection of emission lines and emission-line-like features such as RRC is much easier in the pre-flare phase, when the continuum is almost fully absent in the RGS range (cf. Fig. 2). The higher continuum level during the flare, but also the post-flare, leads to large uncertainties on line properties and often just upper limits on line contribution.

4.2.1. Comparison to previous studies of Vela X-1

We compare the results of our analysis of the time-averaged spectra with the work of Mao et al. (2019), who have performed an automated line search through the entire RGS archive, including this observation of Vela X-1. They analyzed the observation as a whole and found emission lines at 914.98, 1023.09, 1213.14, and 1483.68 eV which we also have detected and identified as Ne IX triplet, Ne X Ly α , Ne X Ly β , and Mg XII Ly α , respectively. Additionally, the authors detected emission lines at 1587.92, 1663.94, and 1733.20 eV, which we have not seen. Their automatic approach did not detect the oxygen lines at 569.10 and

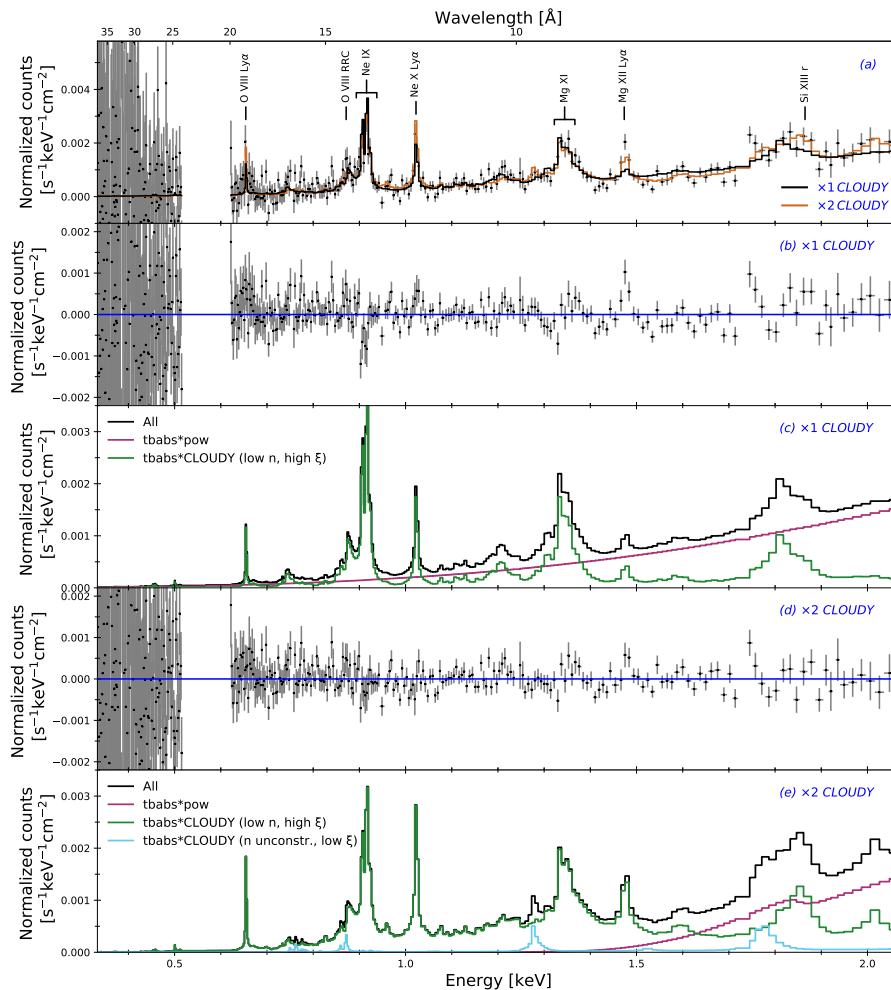


Fig. 7. RGS 2 pre-flare spectrum fitted with photoionized CLOUDY models. Panel (a): the spectrum and the best fit models with one and two CLOUDY components; panel (b): model residuals (one CLOUDY component); panel (c): the total model with one CLOUDY component in black and the additive model components in color; panel (d) and (e): similarly as in panel (b) and (c) but for a model with two CLOUDY components.

654.12 eV and the Mg xi, Mg vii, and Mg vi lines between 1270 and 1360 eV that we detect (cf. Tab. 3). Mao et al. (2019) also observed an emission feature at 824.07 eV that lacks a clear identification.

A comparison between Mao et al. (2019) and our results is not straightforward. The two analyses use different statistical criteria to establish the significance of a line, as well as different assumptions on the underlying continuum. Moreover, the Mao et al. (2019) catalogue does not provide an assessment of the statistical significance of each line, besides the statistical error on the line normalization. It should be born in mind that the Mao et al. (2019) catalogue is the outcome of an automatic, homogeneous analysis on a large number of spectra of completely different classes of objects. Results are not validated on individual observations, because the catalogue is intended to be used for statistical studies only.

As discussed in Sec. 3.3.1, the line velocities we obtain for the detected lines, both in time-averaged and time-resolved spectra, are consistent with zero and thus not reported. The RGS resolution at energies above ~ 1 keV corresponds to line speeds of ~ 2100 – 2400 km/s. For example, line velocities of Si lines in the *Chandra* observations of Vela X-1 in Grinberg et al. (2017) have magnitudes of 1450 km/s or less. Below ~ 0.4 – 0.6 keV, the RGS resolution allows measuring line velocities with magnitudes of ~ 750 – 1500 km/s. However, at lower energies our spectra have a comparatively lower statistical quality.

A plethora of line features has been detected in Vela X-1 previously with different instruments, starting with *ASCA* (Nagase

et al. 1994), but none of the other high resolution observations covers a flare as extreme as the one shown here. To our best knowledge, we also present the first firm detection of oxygen and the first detection of nitrogen transitions in the source. In particular, Sako et al. (1999) included oxygen into their differential emission measure distribution analysis and reported an upper limit detection of O vii as well as a possible detection of O viii RRC. The *Chandra*-HETG analyses (Schulz et al. 2002; Goldstein et al. 2004; Watanabe et al. 2006; Hell et al. 2016; Grinberg et al. 2017) do not cover such low energies. On the other hand, the HETG observations were able to constrain the weak lines from lower ionization states, especially of silicon, during highly absorbed episodes that our analysis lacks the resolution for.

4.2.2. Presence of H- and He-like ions

The detection of H- and He-like ions is indicative of the presence of hot, highly ionized gas. When modelling the Ne ix triplet in all phases, we constrained the centroid energy of the triplet components relative to that of the Neix Ly α transition as dictated by the atomic physics, while for the Mg xi triplet visible in the RGS 2 pre-flare phase, the intercombination and forbidden lines were tied to the resonance line in a similar manner. The supporting assumption implies that these lines have the same Doppler shift and thus belong to the same dynamic system, which seems justified given our data and previous measurements.

As previously mentioned in Sec. 3.4, the $R = f/i$ and $G = (f + i)/r$ diagnostic parameters obtained from line inten-

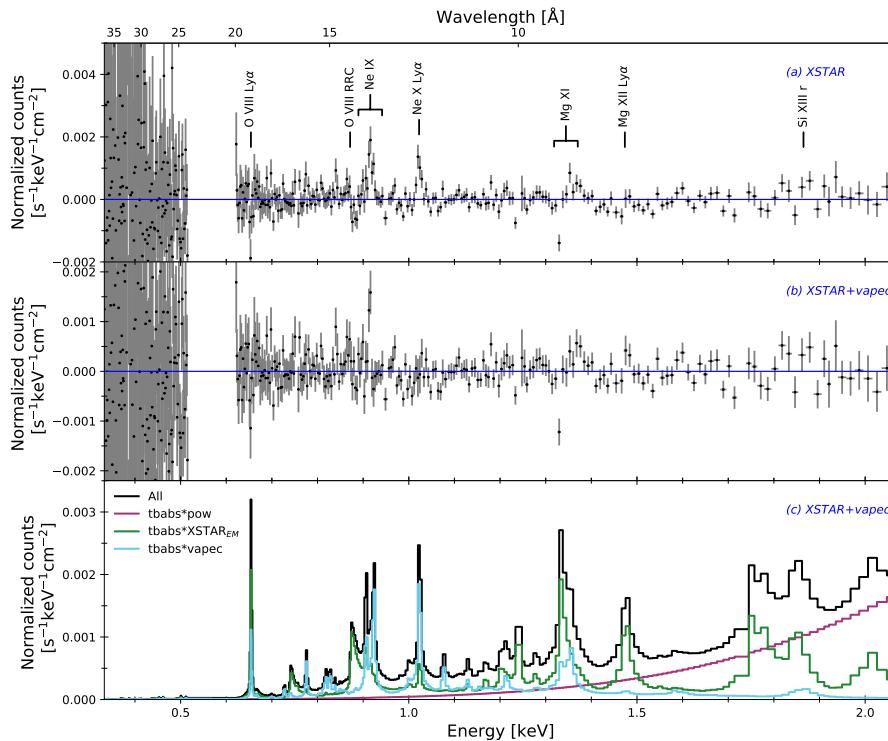


Fig. 8. RGS 2 pre-flare spectrum fitted with a photoionized XSTAR model and a collisionally ionized XSPEC model vapec. Panel (a): model residuals (one XSTAR component); panel (b): model residuals (XSTAR+vapec components); panel (c): the total XSTAR+vapec model in black and the additive model components in color.

sities of He-like atoms can be used as probes for density, temperature, and ionization processes in plasmas. The value of R decreases with increasing electron density, while $G < 4$ in general and $G \sim 1$ in particular indicates the presence of a collisionally ionized component (Porquet & Dubau 2000; Porquet et al. 2001). However, these limits have to be treated with caution since for accretion-powered pulsars the G parameter can reach low values in a photoionized plasma due to resonance scattering that enhances the resonance line (e.g., Schulz et al. 2002; Wojdowski et al. 2003). Additionally, the strong UV field of a B star can lead to reduced strength of the forbidden line as the metastable $1s2s^3S_1$ level gets depopulated (Gabriel & Jordan 1969; Blumenthal et al. 1972; Mewe & Schrijver 1978; Porquet et al. 2001). In particular, in Sect. 4.3.1 we test the importance of the star’s UV contribution by calculating photoionization models without the stellar contribution to the input continuum. Combined with the large measurement errors on R and G in all cases except perhaps the Ne IX in the pre-flare phase (Tab. 7 and Fig. 6), this precludes us from drawing firm conclusions on the state of the plasma from the He-ratios in this observation. The presence of the lines, however, indicates their potential for future missions with higher throughput and similar resolution at the respective triplet-energies, e.g., *XRISM* and *Athena* (see Sect. 4.4).

4.2.3. Low ionization lines

In the flare phase, we also find absorption lines corresponding to the transitions of fluorescent Mg vi K α and Mg vii K α indicating the presence of a less ionized component along the line of sight. Grinberg et al. (2017), who analyzed Vela X-1 at $\phi_{\text{orb}} = 0.25$ with *Chandra*-HETG, detected these transitions during the brighter part of their observation as well. Variable absorption has also been previously detected in the high mass X-ray binary Cygnus X-1 (Hirsch et al. 2019). An explanation for the presence of such less ionized material would be denser, colder regions in the wind, such as wind clumps that are expected to be present in

the wind of O/B giants (e.g., Oskinova et al. 2012; Sundqvist & Owocki 2013) or parts of the accretion structure (e.g., Blondin et al. 1990; Manousakis & Walter 2015; El Mellah et al. 2018).

4.2.4. Line response to changes in the continuum

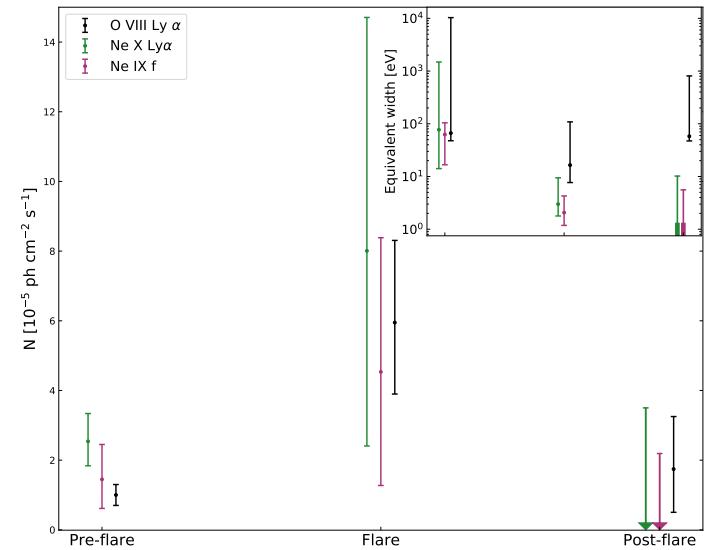


Fig. 9: Main plot: time evolution of the line intensity of the O VIII Ly α , Ne IX f, and Ne X Ly α lines during the pre-, during and post-flare phases with uncertainties (90 % confidence interval for two interesting parameters). **Inset plot:** similarly to above, time evolution of the equivalent width.

To study potential spectral line changes due to the flare, we investigated the line intensity of O VIII Ly α , Ne IX f, and Ne X Ly α in the three phases, as shown in Fig. 9. To do so, we modelled the lines of the Ne IX triplet as three separate Gaussian components

Table 8: Best fit values of the photoionized CLOUDY models in the pre-flare phase (RGS 2).

Model	Parameter	Value
1× CLOUDY		
tbabs	$N_{\text{H}} [10^{22} \text{ cm}^{-2}]$	< 0.2
powerlaw	Γ	$-2.9^{+0.5}_{-0.6}$
cflux	$\log F [\text{erg cm}^{-2}\text{s}^{-1}]$	$-11.62^{+0.05}_{-0.06}$
tbabs	$N_{\text{H}} [10^{22} \text{ cm}^{-2}]$	$1.0^{+0.2}_{-0.1}$
CLOUDY	$\log n_{\text{e}} [\text{cm}^{-3}]$	$8.2^{+0.4}_{-0.2}$
CLOUDY	$\log \xi [\text{erg cm s}^{-1}]$	$3.3^{+0.6}_{-0.1}$
cflux	$\log F [\text{erg cm}^{-2}\text{s}^{-1}]$	$-11.0^{+0.6}_{-0.1}$
C-stat		591
ν		430
2× CLOUDY		
tbabs	$N_{\text{H}} [10^{22} \text{ cm}^{-2}]$	$12.7^{+4.8}_{-3.4}$
powerlaw	Γ	$5.86^{+3.95}_{-6.75}$
cflux	$\log F [\text{erg cm}^{-2}\text{s}^{-1}]$	$-11.81^{+0.07}_{-0.08}$
tbabs	$N_{\text{H}} [10^{22} \text{ cm}^{-2}]$	$0.96^{+0.07}_{-0.09}$
CLOUDY	$\log n_{\text{e}} [\text{cm}^{-3}]$	$8.4^{+0.4}_{-0.4}$
CLOUDY	$\log \xi [\text{erg cm s}^{-1}]$	$3.61^{+0.03}_{-0.04}$
cflux	$\log F [\text{erg cm}^{-2}\text{s}^{-1}]$	$-11.02^{+0.05}_{-0.05}$
tbabs	$N_{\text{H}} [10^{22} \text{ cm}^{-2}]$	$0.96^{+0.07}_{-0.09}$
CLOUDY	$\log n_{\text{e}} [\text{cm}^{-3}]$	unconstrained
CLOUDY	$\log \xi [\text{erg cm s}^{-1}]$	$1.7^{+0.5}_{-0.7}$
cflux	$\log F [\text{erg cm}^{-2}\text{s}^{-1}]$	$-12.3^{+0.3}_{-0.4}$
C-stat		529
ν		426

Notes. N_{H} stands for hydrogen column density, Γ denotes photon index, N is normalization, n_{e} is electron number density, $\log \xi$ denotes ionization parameter, z is red shift, F is flux calculated between 0.33 and 2.1 keV, ‘C-stat’ denotes the C-statistic values, and ν stands for degrees of freedom.

using the build in XSPEC function and not as a triplet (cf. Sec. 3.3.1). We also investigated the Ne IX (r) line in the same manner, but we do not include this result in Fig. 9 because the line intensity in this case was consistent with zero in the flare and post-flare phases. All three lines shown follow the same trend reaching an intensity peak during the flare and reducing afterwards. This trend is consistent with what is seen for the other lines as listed in Tab. 4. The concurrent trend of the intensity of the He- and H-like Ne transitions increasing during the flare could be explained by more material becoming ionized as the irradiation of the wind increases.

In the analysis of the EPIC-pn data of our observation, Martínez-Núñez et al. (2014) saw a peak in Fe K α line flux during the flare, i.e. a similar behavior of the line flux as we observe here for other lines. The measured energy of the Fe K α line is compatible with ionization stages of iron up to Fe XVIII, and thus

Table 9: Same as Tab. 8 but using photoionized XSTAR models and a collisionally ionized XSPEC model vapec.

Model	Parameter	Value
XSTAR		
tbabs	$N_{\text{H}} [10^{22} \text{ cm}^{-2}]$	< 0.2
powerlaw	Γ	$-3.2^{+0.7}_{-0.9}$
cflux	$\log F [\text{erg cm}^{-2}\text{s}^{-1}]$	$-11.74^{+0.06}_{-0.04}$
tbabs	$N_{\text{H}} [10^{22} \text{ cm}^{-2}]$	$1.19^{+0.05}_{-0.04}$
XSTAR	$n [\text{cm}^{-3}]$	$> 2.28 \times 10^{11}$
XSTAR	$\log \xi [\text{erg cm s}^{-1}]$	$2.47^{+0.09}_{-0.07}$
cflux	$\log F [\text{erg cm}^{-2}\text{s}^{-1}]$	$-10.70^{+0.08}_{-0.03}$
C-stat		664
ν		430
XSTAR + vapec		
tbabs	$N_{\text{H}} [10^{22} \text{ cm}^{-2}]$	< 3.2
powerlaw	Γ	$-5.2^{+3.5}_{-1.2}$
cflux	$\log F [\text{erg cm}^{-2}\text{s}^{-1}]$	$-11.73^{+0.07}_{-0.06}$
tbabs	$N_{\text{H}} [10^{22} \text{ cm}^{-2}]$	$1.3^{+0.1}_{-0.1}$
XSTAR	$n [\text{cm}^{-3}]$	$> 1.6 \times 10^{11}$
XSTAR	$\log \xi [\text{erg cm s}^{-1}]$	$2.41^{+0.19}_{-0.06}$
cflux	$\log F [\text{erg cm}^{-2}\text{s}^{-1}]$	$-10.8^{+0.2}_{-0.2}$
tbabs	$N_{\text{H}} [10^{22} \text{ cm}^{-2}]$	$1.3^{+0.1}_{-0.1}$
vapec	$kT_{\text{e}} [\text{keV}]$	$0.36^{+0.04}_{-0.05}$
vapec	$[\text{Fe}/\text{H}]$	< 0.2
cflux	$\log F [\text{erg cm}^{-2}\text{s}^{-1}]$	$-11.07^{+0.09}_{-0.12}$
C-stat		559
ν		427

Notes. N_{H} stands for hydrogen column density, Γ denotes photon index, N is normalization, n is gas number density, $\log \xi$ denotes ionization parameter, ‘C-stat’ denotes the C-statistic values, and ν stands for degrees of freedom.

with iron ions that could be produced in the same plasma that creates the lines we observe with RGS. We, in particular, check that this is indeed the case by looking at the iron ionic abundances corresponding to the best fit photoionization models in Sect. 3.5.

In the inset plot in 9 we also show the evolution of the equivalent width (EW) throughout the flare. We provide these measurements for convenience of comparisons only and in particular emphasize that given the RGS energy range and the strongly varying absorption, the true uncertainties on the continuum and thus on the intrinsic EWs are impossible to estimate with the data we have.

4.2.5. RRC features

RRC features arise as free electrons recombine with ions in a plasma. Since the kinetic energy of the electrons is not quantized, but is determined by the Maxwellian speed distribution, its excess w.r.t. the threshold energy of the atomic level will be transferred to the emitted photon. The intensity of an RRC feature is proportional to the electron density, while its width depends on electron temperature. In an overionized plasma, where the electron temperature is much lower than the ionization temperature, the RRC features are expected to be rather narrow and line-like compared to broad RRC in collisional plasmas.

The detected O VIII RRC at 865.25 eV in the pre-flare phase corresponds to $kT_e = 13.9_{-6.8}^{+17.6}$ eV and $T_e \sim 1.5 \times 10^5$ K. Schulz et al. (2002) analyzed Vela X-1 in eclipse with *Chandra*-HETG and found a similar value, $kT_e \sim 10 \pm 2$ eV and $T_e \sim 1.2 \times 10^5$ K, using the Ne X RRC at 1363.52 eV. In our case, the Ne X RRC feature cannot be resolved from the Mg XI because of RGS's resolution, and we obtain a reasonable fit using three Gaussian lines to represent the He-like Mg-triplet only. If we include the RRC into our model, we only obtain an upper limit $kT_e \leq 12.6$ eV and $T_e \leq 1.5 \times 10^5$ K (Tab. 5). The temperatures we obtain from RRC are typical for photoionized plasmas.

We note that, in principle, we should be able to detect a shift in the ionization balance that would allow us to constrain the density independently by measuring the change in the intensity of the O VIII RRC. Given that the timescale of the flare is $\sim 10^4$ sec (see Fig. 1), and assuming that recombination timescales are of the order $\sim 10^{12}/(n_e \cdot \text{cm}^{-3})$ sec, with the electron density $n_e = 10^8 \text{ cm}^{-3}$ (CLOUDY results in Tab 8), these timescales become comparable. This would be valuable since the density sensitivity of the $n = 1-2$ He-like ion spectra is likely to be compromised by the UV radiation field of the companion star. Unfortunately, the quality of the spectra in this work does not allow us to measure such changes.

4.3. Photoionization modelling

4.3.1. CLOUDY modelling

As previously mentioned in Sec. 3.5.2, one CLOUDY component is sufficient to fit the emission lines in the RGS 2 pre-flare spectrum (see Fig. 7). Nevertheless, two CLOUDY components manage to fit the data better, which is unsurprising given our expectation of a complex structure of the clumpy stellar winds in Vela X-1 (e.g., Sako et al. 1999; Schulz et al. 2002). From the best fit values in Tab. 8, we can conclude that the two CLOUDY components correspond to two different phases of the plasma: the first one being highly ionized ($\log \xi/\text{erg} \cdot \text{cm} \cdot \text{s}^{-1} \approx 3.6$) and the second with a lower ionization ($\log \xi/\text{erg} \cdot \text{cm} \cdot \text{s}^{-1} \approx 1.7$). The highly ionized component is rather thin, with $\log n_e/\text{cm}^{-2} \approx 8$, but our data do not allow to constrain the density of the second, less prominent CLOUDY component. Less ionized components have, however, been interpreted as colder, denser clumps in other works previously, especially such that employed *Chandra*-HETG observations and could resolve lower ionization stages of elements such as Si and Mg (Schulz et al. 2002; Grinberg et al. 2017). The number density of the component with the higher ionization corresponds to the value we calculate for the average number density of the wind at the location of the neutron star in Sec. 3.5.1. If there are indeed two components present and the wind is structured with more and less dense regions, a natural identification would be for the thinner component to be more ionized than the denser.

We remind the reader to not take the values of the photo index in the power laws in Tabs. 8 and 9 literally. As we mentioned in Sect. 3.1, our model of the continuum does not constrain its physical properties.

In addition, we also investigate the effects of the UV radiation on the produced models. To do so, we have created an additional set of CLOUDY models for the same input parameters as above with one exception: the illuminating continuum needed for the simulation now did not contain the blackbody contribution from the supergiant (cf. Sec. 3.5.1). The obtained best fits had a significantly poorer quality than the model with the UV contribution ($\Delta\text{C-statistic} \approx 150$ for the model with two CLOUDY components) and, among others, predicted too strong forbidden lines of the Ne IX and Mg XI triplets. This clearly indicates that UV line pumping is important for such simulations, as expected (cf. also Sec. 4.2.2).

Ions at significantly different ionization stages found together in the spectrum of Vela X-1 in this work (Mg) and other observations, for example, in the *Chandra* observations in Grinberg et al. (2017) (Mg and Si), reveal the complex structure of this system. Describing it as a combination of a small number of discrete photoionization components is likely to be an oversimplification of a rather complex astrophysical system, and more sophisticated plasma models that, for example, can simulate multi-phase plasmas, are needed for these purposes.

4.3.2. XSTAR modelling

In addition to the photoionized CLOUDY models, we also address XSTAR modelling. Comparing results from the two codes is crucial to check for consistency of our results but also to cross-check the codes.

Similarly as for CLOUDY, we allowed two free parameters, the gas density, n , and the logarithm of the ionization parameter, $\log \xi$, when generating photoionized models with XSTAR. However, here n varied between $3 \times 10^7 \leq n/\text{cm}^{-3} \leq 3 \times 10^{11}$ in 10 steps. We used the same values for the turbulent velocity from Sander et al. (2018), assumed solar abundances and spherical geometry of the nebula as well as fixed the column density at 10^{23} cm^{-2} .

To obtain an optimal fit implementing XSTAR models, we have tested different combinations of additive emission and multiplicative absorption XSTAR models as well as combinations of XSTAR with models of optically-thin, collisionally ionized plasma. For the latter component we used *vapec* (Smith et al. 2001), allowing to fit separately and independently the elemental abundances.

In panels (a) and (b) in Fig. 8 we present the results for the two simplest combinations of models that have produced the best results. Panel (a) in Fig. 8 shows the residuals of the first XSTAR model (pre-flare RGS 2 spectrum) which consists of a power law continuum absorbed by *tbabs* with a single additive XSTAR emission model absorbed by a separate *tbabs* component, similarly to the overall CLOUDY model described in Eq. 3. In panel (b) in Fig. 8 we show the residuals of the model with an absorbed continuum as above and an additive XSTAR model together with one *vapec* component, both of them being absorbed by the same *tbabs* (cf. Eq. 3).

It is evident that a single XSTAR component is unable to properly fit the Ne IX triplet and Ne X Ly α line in the pre-flare RGS 2 spectrum. The differences in the fit quality are also visible from the C-statistic values in Tab. 9, and the strong contribution of the *vapec* component at lower energies is clearly shown in panel (c) in Fig. 8. The quality of the fits does not significantly improve if

an additional XSTAR component is added to the model ($\Delta C/\Delta v = 0.1/7$).

The `vapc` component is highly significant ($\Delta C/\Delta v = 105/3$), provided that its iron abundance is low ($Z_{\text{Fe}} < 0.2$). A higher Fe abundance would imply much stronger Fe L emission lines than seen in the data. The best-fit XSTAR component requires a higher gas density ($n > 2 \times 10^{11} \text{ cm}^{-3}$) and an intermediate ionization parameter ($\log \xi \simeq 2.4 \text{ erg} \cdot \text{cm} \cdot \text{s}^{-1}$) with respect to the CLOUDY best-fit components (compare Tab. 8 with Tab. 9). One should also note that the estimated temperature of the collisionally-ionized component $kT_e = 0.36 \text{ keV}$ corresponds to a temperature of the order of 10^6 K , which would imply highly ionized Fe ions, e.g., Fe xxv, Fe xxvi. The Fe K α line from Martínez-Núñez et al. (2014), on the contrary, is consistent with Fe ionized up to Fe xviii.

The difference in the best-fit parameters between CLOUDY- and XSTAR-based fit is puzzling. Observationally, it is driven by the fact that XSTAR models do not properly fit the Ne xi triplet in the Vela X-1 spectrum, in particular the enhanced ratio between the resonant and the forbidden component, as well as the intensity of the Ne x Ly α . Similar issues are apparent in the fit of the Mg triplet. Since the main difference between CLOUDY and XSTAR in the fit of these spectra seems to be focused on the relative importance of the resonant lines in the He-like triplets, it is likely due to the different assumed geometries and treatment of resonant scattering. We do not have any clear evidence for the collisional phase of the wind plasma, which seems needed because XSTAR cannot reproduce the resonant lines at the expenses of the iron abundance, while CLOUDY finds an easier self-consistent solution.

Investigating the origin of these differences remains beyond the scope of this paper. We note that the two codes differ for the treatment of the physical processes underlying the resolution of the radiative transfer equation. The publicly available version of XSTAR, for instance, does not include the treatment of continuum photoexcitation, an issue affecting, for example, the prediction of the Fe L lines in photoionized plasma (Guainazzi et al., in prep.). Previous benchmarks in the framework of AGN outflow studies (Mehdipour et al. 2016) show differences up to 30% in the optical depth of *absorption* lines produced at $\log \xi$ comprised between 1 and 2, and of about 20% at higher ionization parameters.

4.4. Future prospects with XRISM and Athena

The next decade will witness a transformational leap in the capability of high-resolution X-ray spectroscopy with the launch of the Japanese-led X-Ray Spectroscopy and Imaging Mission (XRISM; Tashiro et al. 2018) and of the ESA Advanced Telescope for High-Energy Astrophysics (*Athena*; Nandra et al. 2013) observatories. XRISM will carry a sensitive microcalorimeter detector in the focal plane of a grazing incident telescope (*Resolve/SXS*; Tashiro et al. 2018) with an energy resolution of 7 eV (requirement) over the whole energy range between 0.2 and 12 keV. The *Resolve* instrument aims at bringing to full maturity the promises ushered by the ill-fated *Hitomi* mission, that was lost after six weeks of operations after yielding the first eV-class resolution spectrum of an extended celestial source (the core of the Perseus Cluster; Hitomi Collaboration et al. 2016). As a proxy of what *Resolve* spectroscopic observations could bring, we show in Fig. 10 the 1.2–2.1 keV 20 ks spectrum that the *Hitomi/SXS* would have observed, assuming the best-fit model as in Tab. 8 and an energy resolution degraded to the *Resolve* requirements from the actual *Hitomi/SXS* reso-

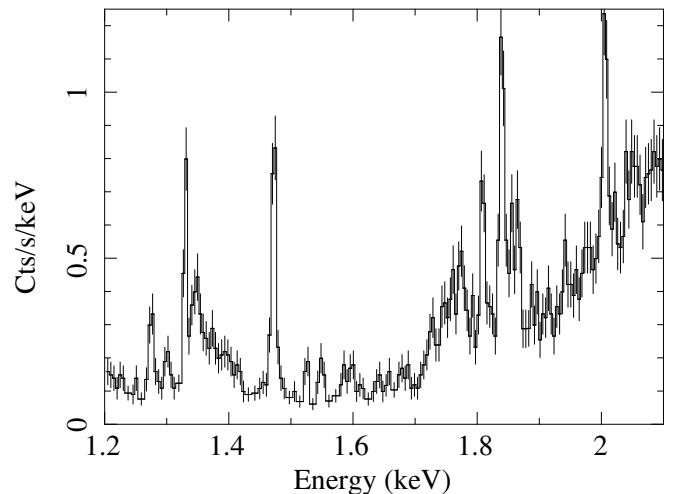


Fig. 10: Simulated spectrum of Vela X-1 assuming the average pre-flare phase best-fit model presented in Tab. 8 and the instrumental response of the *Hitomi/SXS* instrument, the resolution of which was degraded to match the requirement of the XRISM/*Resolve*. Only the energy range between 1.2 and 2.1 keV is shown. The exposure time is 20 ks.

lution of $\sim 5 \text{ eV}$ that outperformed the instrument requirements (Eckart et al. 2018). The expected statistical accuracy of diagnostic ratios based on the He- and H-like lines of Mg and Si are: $\simeq 15\%$ for the Mg xii/Mg xi, $\simeq 20\%$ for the Si xiv/Si xiii, $\simeq 25\%$ for the Mg xi (r)/Mg xi (f), and $\simeq 40\%$ for the Si xiii (r)/Si xiii (f) line ratios, respectively. This would correspond to a statistical error of 35% and 45% on the determination of the R and G parameters for Mg xi, respectively. A 50 ks observation would improve the accuracy to 20% and 30%, respectively. The diagnostic parameters of the Si xiii line would be harder to constrain. One would need at least 30 ks to obtain a statistical accuracy better than 50%. For a 50 ks observation, the estimated statistical accuracy for the Si xiii R and G parameters would be 25% and 35%, respectively.

The full power of phase-resolved spectroscopy will be unfolded with the launch of *Athena* (Nandra et al. 2013), thanks to the unprecedented combination of effective area and spectral resolution (2.5 eV at 7 keV) ensured by the X-Ray Integral Field Unit (X-IFU; Barret et al. 2018). In Fig. 11 we show a simulation of the predicted Vela X-1 X-IFU spectrum corresponding to a net exposure time of 300 s (e.g., commensurate to the Vela X-1 neutron star period). In the simulation, we used the best-fit model as in Tab. 8. We assumed the current mirror configuration corresponding to effective area requirements of 1.4 m^2 at 1 keV, and 0.25 m^2 at 7 keV. Strong H- and He-like lines of Ne, Mg, and Si are visible by-eye, and their intensity and profile can be studied with an accuracy at the percent level.

We note that both for *Athena* and *XRISM* simulations we used the nominal responses despite Vela X-1 being a rather bright source. In the case of *XRISM*, at the average observed flux level ($\sim 5\%$ of a Crab), the *XRISM/Resolve* on-board processor should be able to process the data nominally. At the highest flux level ($\sim 10\%$ of a Crab), the processor could be marginally affected, but this is an unlikely occurrence. Regarding *Athena*, the count rate of Vela X-1 in the pre-flare state is ~ 1900 counts/s, corresponding to about 20 times the counts expected from a mCrab source. At 50 mCrab, for instance, there is a penalty to pay in terms of throughput for the highest resolution events of

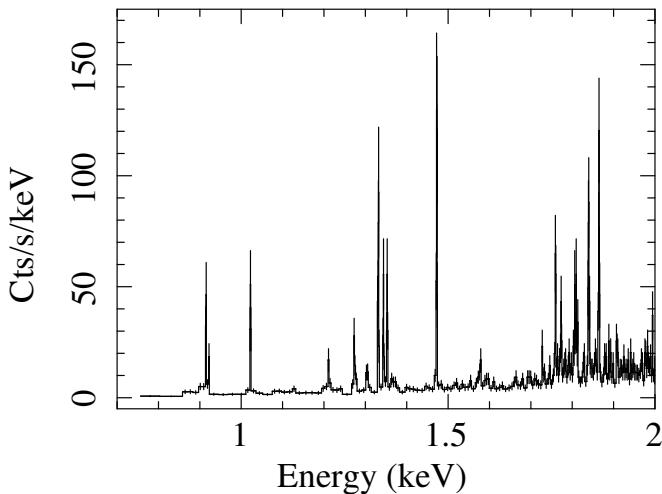


Fig. 11: Simulated *Athena*/X-IFU spectrum of Vela X-1 assuming the average pre-flare phase best-fit model presented in Tab. 8. Only the energy range between 0.8 and 2 keV, is shown. The exposure time is 300 s.

~30% (see Peille et al. 2018, for example). This means that the nominal 300 s above correspond to effective 400 s exposure time.

5. Summary & Outlook

In the current study of the wind accreting HMXB Vela X-1, we investigated the reaction of the clumpy stellar winds to an increase of the intrinsic emission from the neutron star. The time-resolved analysis showed many emission and absorption features, especially prominent in the pre-flare phase against the low, highly absorbed continuum. Because of the strong increase of the continuum contribution during and after the flare, many of these line features disappear or weaken w.r.t. the continuum.

- We detect emission lines corresponding to H- and He-like ions of N, O, Ne, and Mg and Si. To our best knowledge, we present the first direct detection of O and N emission in Vela X-1.
- The detected low ionization of Mg K α absorption lines indicate the presence of colder, less ionized material in the line of sight potentially originating from the clumpy winds or parts of the accretion structure, although the latter is not expected to be strong at the given orbital phase.
- The RRC feature of oxygen yielded a plasma temperature of $T_e \sim 1.5 \times 10^5$ K, which is typical for photoionized plasmas.
- Our attempt to probe the density and temperature of the wind plasma through the R and G diagnostic parameters resulted in large uncertainty ranges, especially in the flare and post-flare phases, while the wind velocity estimations appeared to be consistent with zero.
- The best-fit model with two CLOUDY components calculated for a photoionized plasma suggests the presence of two distinct gas regions with high and low ionization levels. A strong UV emission from the supergiant companion is required to achieve a good agreement between models and data. Spectral modelling with XSTAR delivers the best result in combination with a collisionally ionized XSPEC model `vapc`.

By simulating Vela X-1 spectra as it might be seen by *XRISM* and *Athena*, we show that this source is an ideal target for obser-

vations by the future missions that have the potential to probe plasma properties with a much higher accuracy and precision.

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Table 4: Identified narrow features (emission and absorption lines and RRC) and upper limits in the RGS 1 and RGS 2 spectra.

Ion	E_{ref} [eV]	Phase								
		Pre-flare			Flare			Post-flare		
		E_{obs} [eV]	N	ΔC	E_{obs} [eV]	N	ΔC	E_{obs} [eV]	N	ΔC
			$\text{[ph cm}^{-2} \text{s}^{-1}\text{]}$	$\times 10^{-5}$		$\text{[ph cm}^{-2} \text{s}^{-1}\text{]}$	$\times 10^{-5}$		$\text{[ph cm}^{-2} \text{s}^{-1}\text{]}$	$\times 10^{-5}$
N vii Ly α	500.71 ^a	499.3	< 0.3	—	499.3 ^{+1.1} _{-1.9}	1.5 ^{+1.6} _{-1.2}	7.3	499.3	< 1.7	—
O vii He- α (i)	569.10 ^b	—	—	—	568.5 ^{+2.4} _{-1.2}	3.0 ^{+2.7} _{-2.1}	9.8	568.52	< 1.4	—
O viii Ly α_{RGS1}	654.12 ^a	—	—	—	653.6 ^{+1.1} _{-1.1}	3.0 ^{+2.7} _{-2.3}	7.6	653.6 ^{+6.5} _{-2.1}	1.7 ^{+1.4} _{-1.2}	6.5
O viii Ly α_{RGS2}	654.12 ^a	654.15 ^{+0.8} _{-1.3}	1.0 ^{+0.3} _{-0.3}	18.8	653.1 ^{+0.9} _{-0.8}	5.9 ^{+2.3} _{-2.1}	58.4	661.1 ^{+2.4} _{-1.7}	1.7 ^{+1.5} _{-1.2}	6.0
O vii He- β	666.08 \pm 0.01 ^c	668.48	< 0.7	—	668.6 ^{+4.7} _{-1.7}	2.3 ^{+1.9} _{-1.6}	6.3	E _{Ovii} He γ - 32.20	< 1.7	—
O vii He- γ	698.32 \pm 0.03 ^c	698.53	< 0.7	—	E _{Ovii} He β + 32.20	< 2.4	—	697.6 ^{+2.8} _{-1.9}	1.5 ^{+1.5} _{-1.3}	6.6
Ne ix (f)	905.74 ^d	$E_{\text{Ne x Ly}\alpha} - 116.89$	1.3 ^{+0.3} _{-0.3}	137.4	$E_{\text{Ne x Ly}\alpha} - 116.89$	5.2 ^{+2.9} _{-3.1}	39.6	$E_{\text{Ne x Ly}\alpha} - 116.89$	0.2 ^{+0.1} _{-0.9}	18.4
Ne ix (i)	915.58 ^d	$E_{\text{Ne x Ly}\alpha} - 106.86$	—	—	$E_{\text{Ne x Ly}\alpha} - 106.86$	—	—	$E_{\text{Ne x Ly}\alpha} - 106.86$	—	—
Ne ix (r)	922.69 ^d	$E_{\text{Ne x Ly}\alpha} - 99.85$	—	—	$E_{\text{Ne x Ly}\alpha} - 99.85$	—	—	$E_{\text{Ne x Ly}\alpha} - 99.85$	—	—
Ne x Ly α	1022.60 ^a	1022.0 ^{+0.9} _{-0.5}	2.5 ^{+0.8} _{-0.7}	68.3	1022.4 ^{+1.5} _{-1.8}	8.0 ^{+6.7} _{-5.6}	11.0	1022.5	< 3.5	—
Ne x Ly β	1211.85 ^a	1212.9 ^{+4.2} _{-6.2}	1.7 ^{+0.9} _{-0.7}	24.6	1212.9	< 10.8	—	1212.9	< 9.8	—
Mg vii K α	1276.81 ^e	1275.1	> -0.3	—	1275.1 ^{+4.9} _{-4.8}	-20.0 ^{+11.1} _{-9.9}	17.0	1275.1	> -6.5	—
Mg vii K α	1289.95 ^e	1293.1	> -0.4	—	1293.1 ^{+8.1} _{-4.5}	-24.5 ^{+10.6} _{-11.4}	24.1	1293.1	> -7.8	—
Mg xi (f)	1332.11 ^d	$E_{\text{Mg xi r}} - 21.15$	1.4 ^{+0.5} _{-0.9}	104.4	$E_{\text{Mg xi r}} - 21.15$	< 6.6	—	$E_{\text{Mg xi r}} - 21.15$	< 0.3	—
Mg xi (i)	1344.33 ^d	$E_{\text{Mg xi r}} - 8.93$	—	—	—	—	—	—	—	—
Mg xi (r)	1353.26 ^d	1350.2 ^{+3.9} _{-3.4}	—	—	—	—	—	—	—	—
Mg xii Ly α	1473.46 \pm 7.0 \times 10 ^{-3a}	1475.4 ^{+6.7} _{-3.0}	3.6 ^{+1.3} _{-1.1}	52.1	1475.4	< 35.3	—	1475.35	< 14.6	—

Table 4: continued.

Ion	E_{ref} [eV]	Phase							
		Pre-flare			Flare			Post-flare	
		E_{obs} [eV]	N	ΔC	E_{obs} [eV]	N	ΔC	E_{obs} [eV]	N
			$[\text{ph cm}^{-2} \text{s}^{-1}]$		$[\text{ph cm}^{-2} \text{s}^{-1}]$			$[\text{ph cm}^{-2} \text{s}^{-1}]$	
			$\times 10^{-5}$		$\times 10^{-5}$		$\times 10^{-5}$	$\times 10^{-5}$	
Si xiii (r)	1864.84 ± 0.05^f	$1856.7^{+23.5}_{-1.7}$	$3.4^{+1.4}_{-1.4}$	6.8	$1869.1^{+0.9}_{-0.6}$	$71.9^{+17.6}_{-17.6}$	15.2	1864.84	< 23.7
									–

Notes. E_{ref} denotes reference energy, E_{obs} is detected energy, N is normalization, and ΔC is the difference in the C-statistic with and without the line in question. The error bars were calculated for a 90% confidence interval and two interesting parameters.

References. ^(a) Erickson (1977) ^(b) Grant et al. (1980) ^(c) Engstrom & Litzen (1995) ^(d) Drake (1988) ^(e) From Behar & Netzer (2002); value for Mg vii K α averages over the two strong transitions listed there ^(f) Hell et al. (2016) with an additional systematic uncertainty of 0.13 eV.

Table 6: Unidentified narrow features, potentially Fe lines, in the RGS 1 and RGS 2 spectra.

Ion	E_{ref} [eV]	Pre-flare			Flare			Post-flare		
		E_{obs} [eV]	N [ph cm $^{-2}$ s $^{-1}$] $\times 10^{-5}$	ΔC	E_{obs} [eV]	N [ph cm $^{-2}$ s $^{-1}$] $\times 10^{-5}$	ΔC	E_{obs} [eV]	N [ph cm $^{-2}$ s $^{-1}$] $\times 10^{-5}$	ΔC
RGS 2										
Fe L XXII	1053.392 ^a	1047.6 ^{+4.0} _{-4.4}	-0.5 ^{+0.4} _{-0.7}	6.2	1056.3 ^{+11.5} _{-26.4}	-5.9 ^{+5.6} _{-5.0}	6.0	1053.4	>-4.3	-
Fe L XXIII	1056.443 ^a	-	-	-	-	-	-	-	-	-
Fe L XVIII	- ^b	1067.6 ^{+4.1} _{-6.3}	-0.5 ^{+0.5} _{-1.0}	4.3	1067.6	>-5.1	-	1067.6	>-2.2	-
Fe L XXII	- ^b	-	-	-	-	-	-	-	-	-
Fe L XVIII	1059.513 ^a	-	-	-	-	-	-	-	-	-
Fe L XVIII	1082.075 ^a	-	-	-	-	-	-	-	-	-
Fe L XXIV	1084.536 ^a	-	-	-	-	-	-	-	-	-
Fe L XXIV	1109.379 ^a	1115.8	< 0.7	-	1115.8 ^{+6.0} _{-12.0}	6.1 ^{+8.7} _{-5.8}	5.0	1115.8	< 6.8	-

Notes. E_{ref} denotes reference energy, E_{obs} is detected energy, N is the normalization, and ΔC is the the difference in the C-statistic with and without the line in question. The error bars were calculated for a 90% confidence interval and two interesting parameters.

References. ^(a) Brown et al. (2002) ^(b) Many possible transition are applicable. See Brown et al. (2002) for details.